

Euclid Quick Data Release (Q1)

The active galaxies of *Euclid*

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ABSTRACT

We present a catalogue of candidate active galactic nuclei (AGN) in the *Euclid* Quick Release (Q1) fields. For each *Euclid* source we collect multi-wavelength photometry and spectroscopy information from Galaxy Evolution Explorer (GALEX), *Gaia*, Dark Energy Survey (DES), Wide-field Infrared Survey Explorer (WISE), *Spitzer*, Dark Energy Survey (DESI), and Sloan Digital Sky Survey (SDSS), including spectroscopic redshift from public compilations. We investigate the AGN contents of the Q1 fields by applying selection criteria using *Euclid* colours and WISE-AllWISE cuts finding respectively 292 222 and 65 131 candidates. We also create a high-purity QSO catalogue based on *Gaia* DR3 information containing 1971 candidates. Furthermore, we utilise the collected spectroscopic information from DESI to perform broad-line and narrow-line AGN selections, leading to a total of 4392 AGN candidates in the Q1 field. We investigate and refine the Q1 probabilistic random forest QSO population, selecting a total of 180 666 candidates. Additionally, we perform SED fitting on a subset of sources with available z_{spec} , and by utilizing the derived AGN fraction, we identify a total of 7766 AGN candidates. We discuss purity and completeness of the selections and define two new colour selection criteria ($JH_{-I_E}Y$ and I_EH_{-gz}) to improve on purity, finding 313 714 and 267 513 candidates respectively in the Q1 data. We find a total of 229 779 AGN candidates equivalent to an AGN surface density of 3641 deg^{-2} for $18 < I_E \leq 24.5$, and a subsample of 30 422 candidates corresponding to an AGN surface density of 482 deg^{-2} when limiting the depth to $18 < I_E \leq 22$. The surface density of AGN recovered from this work is in line with predictions based on the AGN X-ray luminosity functions.

Key words. Galaxies: active, Catalogues, Surveys

1. Introduction

Active galactic nuclei (AGN) are some of the most powerful sources in the Universe. With bolometric luminosities up to $L_{\text{bol}} = 10^{48} \text{ erg s}^{-1}$ (Padovani et al. 2017), these objects exist at the centres of massive galaxies and emit immense amounts of non-stellar radiation (Peterson 1997; Netzer 2015; Alexander & Hickox 2012; Combes 2021) due to the accretion of matter onto a super-massive black hole (SMBH) and its surrounding accretion disc embedded in a dusty, clumpy, obscuring torus (Shakura & Sunyaev 1973; Zel’dovich & Novikov 1964; Rees 1984; Peterson 1997; Antonucci 1993; Netzer 2015).

The activity of a SMBH is closely related to the properties of its host galaxy through energetic winds providing valuable feedback, which gives rise to relationships such as the $M-\sigma$ relation (Silk & Rees 1998; Merritt 2000; Haehnelt & Kauffmann 2000; Sahu et al. 2019), the black hole mass to bulge mass relation (Magorrian et al. 1998; Häring & Rix 2004; Kormendy & Ho 2013), and even the black hole mass to host galaxy stellar mass relation (Bandara et al. 2009; Shankar et al. 2016), indicating that understanding the many types of AGN is key to deciphering the origin and evolution of galaxies. This is why identifying AGN in their different states of accretion and obscuration is fundamental to build a full picture of the evolution and properties of their host galaxies (Harrison & Ramos Almeida 2024).

Our current census of AGN is incomplete, partly because we lack a universal diagnostic tool to identify the overall population of these objects (Lacy et al. 2004; Stern et al. 2005,

2012; Kirkpatrick et al. 2012), leading to samples whose properties are strongly biased by their selection methods (Hickox & Masini 2019; Cann et al. 2019; Hviding et al. 2024). AGN diagnostics have been developed for most wavelength ranges. Some of the most common techniques involve using radio observations (Mushotzky 2004; Smolčić et al. 2017; Hickox & Alexander 2018), X-ray emission (Pounds 1979; Brandt & Alexander 2015; Lusso & Risaliti 2016), emission line diagnostics (Baldwin et al. 1981; Veilleux & Osterbrock 1987; Osmer & Hewitt 1991; Greene & Ho 2005), variability diagnostics (Ulrich et al. 1997; Kawaguchi et al. 1998; Paolillo et al. 2004), or colour criteria (Sandage 1971; Koo & Kron 1988; Richards et al. 2001; Stern et al. 2005; Wang et al. 2016), as well as machine-learning methods (Fotopoulou & Paltani 2018). However, all of these techniques have their own limitations. For instance, AGN selection in the ultra-violet (UV), optical, and soft X-rays, are affected by dust and gas obscuration, creating a bias against obscured AGN, also known as Type II AGN (Gilli et al. 2007; Treister et al. 2009; Bornancini et al. 2022).

The infrared (IR) regime is a powerful alternative for AGN identification, particularly for obscured sources (Hickox & Alexander 2018; Bornancini et al. 2022; Calabrò et al. 2023). IR radiation is created by UV and optical accretion disc photons that are absorbed by a surrounding dusty torus and re-emitted in the IR (Antonucci 1993; Urry & Padovani 1995; Mor & Netzer 2012). This theoretically means that by using IR diagnostics, one should be able to detect a sizeable population of obscured AGN (Calabrò et al. 2023). Previous works have already developed both spectroscopic and photometric selection approaches

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for IR surveys (de Grijp et al. 1987; Clavel et al. 2000; Lacy et al. 2004; Stern et al. 2005, 2012; Assef et al. 2013). In particular, colour criteria present a rapid and inexpensive approach for cataloguing these sources. Nonetheless, these techniques still have limitations, and not only are some AGN types still missed, but contaminants also play a role in this selection regime (Borncini et al. 2022).

Euclid is an optical and near-IR (NIR) European Space Agency (ESA) mission (Laureijs et al. 2011). The details of the instruments and its scientific goals can be found in Euclid Collaboration: Mellier et al. (2024). Briefly, *Euclid* will observe approximately 14 000 deg² of the extra-galactic sky while undertaking two surveys during its expected 6-year lifetime. The Euclid Wide Survey (EWS, Euclid Collaboration: Scaramella et al. 2022), which will observe about 14 000 deg² with a visible depth of $I_E = 26.2$, and the Euclid Deep Survey (EDS), which will concentrate on three different areas of the sky covering over 53 deg² with a visible depth of $I_E = 28.2$ (Euclid Collaboration: Mellier et al. 2024). It is expected that *Euclid* will be able to detect billions of sources, of which at least 10 million are anticipated to be AGN identified through a combination of its Visible Camera (VIS, Euclid Collaboration: Cropper et al. 2024) and Near-Infrared Spectrometer and Photometer (NISIP, Euclid Collaboration: Jahnke et al. 2024) instruments (Euclid Collaboration: Selwood et al. 2025; Euclid Collaboration: Bisigello et al. 2024; Euclid Collaboration: Lusso et al. 2024). This will increase our known number of AGN dramatically, meaning that – with the right target selection tools and strategic overlap with other multi-wavelength data sets – *Euclid* will play a crucial role in creating a more complete AGN census.

For this reason, and in anticipation of *Euclid*'s first Quick Data Release (Euclid Quick Release Q1 2025), which constitutes a first visit to the Euclid Deep Fields (EDFs), covering a total area of 63.1 deg², various approaches to identify AGN have already been developed. In particular, Euclid Collaboration: Bisigello et al. (2024) carried out a systematic study to find the best colour-selection criteria for AGN based on *Euclid*'s photometry. That paper emphasizes distinct selection methods that may be more appropriate for either the EWS or EDS. Although the purity of these diagnostics could be enhanced, they provide an excellent foundation for examining the populations present in the Q1 data.

The Q1 data provides observations of the EDF-North (EDF-N), EDF-South (EDF-S), and EDF-Fornax (EDF-F), at the depth of the EWS (Euclid Collaboration: Aussel et al. 2025). The three EDF regions were selected primarily due to the nearly perennial visibility of ecliptic poles under the survey strategy (Euclid Collaboration: Mellier et al. 2024). The overlap of the EDF regions with multi-wavelength external surveys provides an excellent opportunity to investigate the multi-wavelength properties of *Euclid*'s sources. Detailed catalogues of detected sources have been produced for various different missions, and collectively, over 20 million AGN candidates have been identified (Assef et al. 2013, 2018; Storey-Fisher et al. 2024; Fu et al. 2024). Studies based on the combination of external catalogues with those created from the *Euclid* data sets will play a crucial role in advancing our understanding of AGN demography and evolution.

In this paper, we present a multi-wavelength AGN candidate catalogue derived from *Euclid*'s photometry in combination with external surveys. In Sect. 2, we introduce and describe *Euclid*'s Q1 source catalogues, along with the external photometric and spectroscopic catalogues utilised in this work. Additionally, we explain how we perform counterpart (CTP) associations for each survey and provide the number of matches found. In Sect. 3,

we explore the various source populations identified in the data, with a primary focus on stellar and AGN candidates. For the AGN candidates, we examine multiple selection methods, both spectroscopic and photometric, and we compare these diagnostic techniques to those used in other Q1 papers (Euclid Collaboration: Roster et al. 2025; Euclid Collaboration: Margalef-Bentabol et al. 2025; Euclid Collaboration: Stevens et al. 2025; Euclid Collaboration: La Marca et al. 2025, Euclid Collaboration: Laloux et al., in prep). Section 4 examines the AGN candidates obtained, compares them with expected results from the literature, and analyses the different AGN populations identified in this work. Finally, in Sect. 5, the overall multi-wavelength AGN catalogue is presented. With Fig. 1 we provide a diagram that illustrates the procedures we follow to compile the AGN candidate catalogue and we indicate the relevant sections of the paper associated with each step throughout. We adopt a Λ CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$ and $\Omega_\Lambda = 0.7$. All magnitudes are in the AB system (Oke & Gunn 1983) unless stated otherwise.

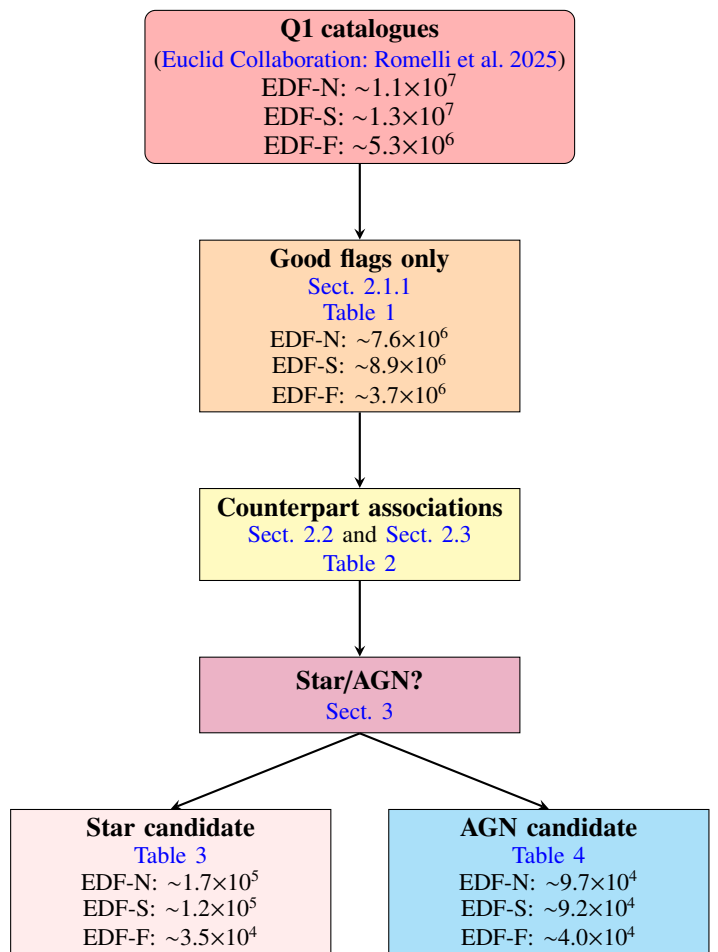


Fig. 1. Sketch outlining the steps adopted in this work to attain the AGN candidate catalogue. We report the number of stellar and AGN candidates for the magnitude range $18 < I_E \leq 24.5$.

2. Data and counterpart associations

This section provides an overview of the data utilised in this study and the nearest-neighbour matching we perform to identify *Euclid*'s multi-wavelength counterparts. We begin with a concise summary of the Q1 catalogues, which serve as the founda-

tion for our AGN catalogue. Subsequently, we split the external data into photometry and spectroscopy, detailing the respective instruments and surveys used.

2.1. *Euclid* Q1 data

Euclid is scheduled to have three major data releases (DRs) over its 6-year nominal mission duration. Detailed descriptions of these releases, as well as information about the mission, are available in [Euclid Collaboration: Mellier et al. \(2024\)](#). However, in addition to the primary releases, there are also interspersed quick releases of smaller volume planned between them. The first of these, the Q1, signifies the initial public release to the scientific community. Details on the data available for the Q1 release can be found in [Euclid Collaboration: Aussel et al. \(2025\)](#), [Euclid Collaboration: McCracken et al. \(2025\)](#), [Euclid Collaboration: Polenta et al. \(2025\)](#), and [Euclid Collaboration: Altieri et al.](#), in prep.

Q1 encompasses a range of data products. Of particular significance for this study are the photometric catalogues generated by the *Euclid* MERge Processing Function (MER, [Euclid Collaboration: Romelli et al. \(2025\)](#)), which include aperture flux measurements with the corresponding errors, quality flags, and morphological information, as well as template fit and Sérsic fit fluxes in each band for all sources detected in the EDFs. Moreover, Q1 includes imaging ([Euclid Collaboration: McCracken et al. 2025](#); [Euclid Collaboration: Polenta et al. 2025](#)) and spectroscopic data ([Euclid Collaboration: Copin et al. 2025](#); [Euclid Collaboration: Le Brun et al. 2025](#)), as well as physical parameter estimations ([Euclid Collaboration: Tucci et al. 2025](#)).

All EDFs have been observed by the four *Euclid* photometric bands, i.e., I_E from VIS in the visible ([Euclid Collaboration: Cropper et al. 2024](#)), and Y_E , J_E , and H_E from NISP in the NIR ([Euclid Collaboration: Schirmer et al. 2022](#); [Euclid Collaboration: Jahnke et al. 2024](#)). These measurements are accompanied by ground-based optical photometry taken with the *ugriz* bands of various instruments, including the Ultraviolet Near-Infrared Optical Northern Survey (UNIONS, Gwyn et al. in prep.) and the Dark Energy Survey ([Abbott et al. 2018](#)), that are re-processed through the official *Euclid* pipelines and homogenised by MER. The breakdown of the photometry available for each EDF and their corresponding instrument can be found in [Euclid Collaboration: Aussel et al. \(2025\)](#).

Additional to the photometric information provided by the *Euclid* catalogues, Q1 also provides spectroscopic catalogues. These data are obtained from NISP-S observations in two red grisms (RGS000 and RGS180) covering the 1206–1892 nm wavelength range. The data reduction process, spectral extractions and data specifics for Q1 spectroscopy are described in [Euclid Collaboration: Copin et al. \(2025\)](#) and [Euclid Collaboration: Le Brun et al. \(2025\)](#). In this work, we focus on the *Euclid* MER photometric catalogues to investigate the source populations present in the Q1 data.

2.1.1. Quality flags and data cleaning

The Q1 photometric catalogues include a number of artefacts, easily identified through a series of flags that are provided as data models. For instance, the reference photometric measurement of a source is given by the FLUX_DETECTION_TOTAL column, and the reliability of this measurement can be assessed using the binary DET_QUALITY_FLAG column. With this flag, a source can be identified to be contaminated by close neighbours,

bad pixels, blending with other sources, saturation, being close to a CCD border, being within the VIS or NIR bright star masks, being within an extended object area, or being skipped by the deblending algorithm. The DET_QUALITY_FLAG is the most informative flag we use to clean the data from the contaminants listed above. Nevertheless, several other flags can also be used to detect contamination in specific bands (i.e., using <band>_FLAG) or contamination by spurious sources (SPURIOUS_FLAG).

In the process of constructing our AGN catalogue from the existing MER Q1 catalogues, we retain only those sources that meet our ‘good flags’ criteria,

- SPURIOUS_FLAG = 0 ,
- <band>_FLAG = 0 ,
- DET_QUALITY_FLAG = 0 | 2 | 512 ,

where DET_QUALITY_FLAG values of 0, 2, and 512 indicate no problems with the data, sources blended together, and sources within an extended object area, respectively.

By applying this ‘good flags’ method, we exclude approximately 32% of the data, resulting in what we from now on refer to as the ‘quality-filtered’ catalogues. Furthermore, considering the varying magnitude limits of the external catalogues we use, we refine the data by dividing them into three magnitude bins: $18 < I_E \leq 21$, $21 < I_E \leq 22$, and $22 < I_E < 24.5$. A detailed breakdown of the number of sources left after these cleaning steps and splitting of the data is provided in [Table 1](#).

2.2. Photometry

In the following sections we discuss the multi-wavelength photometric data used to identify the different source populations, and the counterpart associations performed in this work, ordered by descending energy. Positional matches with the external surveys were performed using the STIL Tool Set (*STILTS* version 3.5-1, [Taylor 2006](#)), which is a package for command-line processing of tabular data, such as astronomical tables. The matches for the three Q1 fields were tailored to account for their different survey coverages. [Table 2](#) indicates the data sets matched to each EDF, the numbers of sources per data set that fall within the Q1 fields, and the number of counterparts found for the quality-filtered versions of the Q1 catalogues.

2.2.1. Ultra-Violet

The UV regime offers insights into some of the most active processes in the Universe that are not observable with optical bands. In this energy range, the EDFs overlap with NASA’s Galaxy Evolution Explorer (GALEX, [Bianchi & GALEX Team 1999](#)). GALEX imaged the sky in two ultraviolet bands: the far-UV (FUV, $\lambda_{\text{eff}} = 1528 \text{ \AA}$); and the near-UV (NUV, $\lambda_{\text{eff}} \sim 2310 \text{ \AA}$). It provided the first UV sky surveys using two observing modes: direct imaging and grism field spectroscopy. GALEX achieved an image full width half maximum (FWHM) of $4'2$ in the FUV and $5'3$ in the NUV. The GALEX GR6/7 data release ([Bianchi et al. 2017](#)) includes millions of source measurements, mostly from the All-Sky Imaging Survey (AIS), with a 5σ limiting magnitude of about 20 in FUV and ~ 21 in NUV. In this work, we use the combined photoobj catalogue, which includes all GALEX programmes: AIS, Medium Imaging survey (MIS), and the Deep Imaging Survey (DIS). GALEX has a lower angular resolution compared to *Euclid*’s VIS point spread function (PSF) FWHM of $0'13$ ([Euclid Collaboration: Mellier et al. 2024](#)). Therefore, when matching between *Euclid* and GALEX, we set the fixed error radius in STILTS to a conservative value of $1'5$. This allows

Table 1. Number of sources in each EDF and impact of quality cuts.

Field	Q1 Catalogue	Good flags ^a	Magnitude bins		
			$18 < I_E \leq 21$	$21 < I_E \leq 22$	$22 < I_E \leq 24.5$
EDF-N	11 378 352	7 573 476	115 606	192 722	2 568 146
EDF-S	13 060 965	8 913 816	127 649	217 375	3 056 362
EDF-F	5 328 489	3 705 597	51 294	88 629	1 326 915

Notes. ^(a) These are the ‘quality-filtered’ catalogues.

Table 2. Surveys cross-matched per *Euclid* field.

Field	Data set	Sources in field	No. matches	Reference
EDF-N	GALEX	225 685	52 663	This work
	<i>Gaia</i>	192 109	43 253	RE25
	WISE-AllWISE	487 397	266 029	This work
	<i>Spitzer</i>	11 378 352	7 573 476	BL25
	DESI	110 459	24 922	This work
	SDSS	326	18	This work
EDF-S	GALEX	225 268	58 093	This work
	<i>Gaia</i>	130 647	35 263	RE25
	DES	4 258 555	3 197 960	This work
	WISE-AllWISE	352 135	268 281	This work
	<i>Spitzer</i>	11 378 352	8 913 816	BL25
EDF-F	GALEX	778 194	230 466	This work
	<i>Gaia</i>	46 500	14 682	RE25
	DES	4 258 555	1 330 109	This work
	WISE-AllWISE	232 079	147 389	This work
	<i>Spitzer</i>	11 378 352	3 705 597	BL25

Notes. The reported matches are between the external survey sources and the quality-filtered *Euclid* catalogues (i.e., only good quality flags).

for a more flexible matching, which in itself is important because sources that might have appeared as blended for the GALEX survey, can potentially be disentangled with *Euclid*’s resolution. The total number of matches we obtain between GALEX and the quality-filtered catalogues is 341 222 (see Table 2 for the breakdown of matches per EDF).

2.2.2. Optical

Historically, optical surveys have been significant for identifying and cataloguing a vast number of sources, therefore enhancing our knowledge of the Universe and the populations found within it. In this energy range, all three EDFs overlap with ESA’s *Gaia* mission (Gaia Collaboration: Prusti et al. 2016). *Gaia*, with its PSF FWHM of 0’1, aims to measure the three-dimensional spatial and the three-dimensional velocity distribution of stars in order to map and understand the formation, structure, and evolution of our Galaxy. The most recent *Gaia* Data Release 3 (DR3, Gaia Collaboration: Vallenari et al. 2023) provides comprehensive source lists that include celestial positions, proper motions, parallaxes, and broadband photometry in the G , G_{BP} (330–680 nm), and G_{RP} (630–1050 nm) passbands, with a limiting depth of $G \approx 21$. Additionally, it offers astrophysical parameters and source class probabilities, including stars, galaxies, and quasars (QSOs) over the entire sky. The Q1 catalogues include *Gaia* IDs from matches performed within the *Euclid* pipeline, which are released as part of the overall Q1 products. The matching performed between these two surveys is explained in Euclid Collaboration: Romelli et al. (2025), hereafter referred to as RE25. The number of identified matches between the quality-

filtered *Euclid* catalogues and *Gaia* is 93 198, the breakdown of which is reported in Table 2.

Moreover, the EDF-S and EDF-F share coverage with the Dark Energy Survey (DES, The Dark Energy Survey Collaboration 2005), which is a ground-based visible and near-infrared imaging survey, aiming to cover 5000 deg² of the southern high Galactic latitude sky. The second DES large data release (DR2, Abbott et al. 2021) contains co-added images and source catalogues, as well as calibrated single-epoch CCD images, from the processing of all six years of DES wide-area survey observations in five broad photometric bands, $grizY$ (Kessler et al. 2015) and all five years of DES supernova survey observations in the $griz$ bands (Diehl et al. 2019), with a detection limit of $g < 25$ and a PSF FWHM typically around 0’8 (Abbott et al. 2021). To perform the counterpart associations between *Euclid* and DES, after investigating different fixed error radii based on the PSF FWHM of both surveys, we set this parameter to 0’55 and obtain a total of 4 528 069 matches (see Table 2 for the breakdown of matches per EDF).

2.2.3. Infrared

The infrared regime provides valuable insights into regions of the Universe that are obscured by dust. In this energy range, the Q1 photometry can be combined with various surveys to enhance our understanding of obscured sources.

In the mid-infrared (MIR), the Wide-field Infrared Survey Explorer (WISE, Wright et al. 2010) is a survey mapping the entire sky in four infrared bands (i.e., W1, W2, W3, W4) centred at 3.4, 4.6, 12, and 22 μ m. Its AllWISE programme combined data from the WISE cryogenic and NEOWISE post-cryogenic sur-

vey (Mainzer et al. 2011) to form the most comprehensive view of the full mid-infrared sky currently available. The AllWISE Data Release, mapping the entire sky and therefore including all three EDFs, provides images with a pixel scale of $1''.375$, source catalogues, multi-epoch photometry catalogues, and reject catalogues up to a detection limit of $W1 < 17.1$. Similar to GALEX, the WISE-AllWISE resolution is not as powerful as that of *Euclid*, having a PSF with FWHM of $6''.1$, $6''.8$, $7''.4$, and $12''$, for its four W1, W2, W3, and W4 bands (Wright et al. 2010). Therefore, when matching between the two surveys, we decided to set the fixed error radius to a conservative value of $1''.5$, allowing for a more flexible matching and resulting in a total number of 681 699 matches, the breakdown of which is reported in Table 2.

To further complement the *Euclid* catalogues, Euclid Collaboration: Bisigello et al. (2025), from now on referred to as BL25, performed forced photometry on Spitzer IRAC images at the position of the *Euclid* sources (i.e., fixed positions). Briefly, starting from the public images by Euclid Collaboration: Moneti et al. (2022) in all four IRAC bands, which include the [3.6], [4.5], [5.6], and [8.0] filters, they first remove the sky background, using a 3×3 pixel filter. Then, the extraction is performed using the position of all *Euclid* sources, both VIS- and NISP- detected, using an aperture with $1''$ radius, resulting in IRAC aperture photometry, which they correct to total, for every *Euclid* source, therefore making the counterpart association unnecessary.

2.3. Spectroscopy

The following sections discuss the spectroscopic data used to identify the different source populations, and the counterpart associations performed in this work.

2.3.1. DESI EDR

The Dark Energy Spectroscopic Instrument (DESI) is a multi-object fibre spectrograph installed at the Mayall-Telescope at Kitt Peak (DESI Collaboration et al. 2022), capable of covering a $3^\circ.2$ wide field of view (Silber et al. 2023). In preparation for its ambitious main survey, a set of survey validation projects (DESI Collaboration et al. 2024a) were conducted with DESI to optimise the final target selection and explore the capabilities and limits of the instrument. These tests consisted of a commissioning data set and a series of survey validations (SV) 1, 2, and 3 (Alexander et al. 2023; Brodzeller et al. 2023; Guy et al. 2023; Lan et al. 2023). The data collected during these pilot surveys were published as the early data release (EDR) of DESI (DESI Collaboration et al. 2024b). It contains spectra of 2 847 435 unique pointings (including sky), which yield 1 202 846 reliable extragalactic spectroscopic redshifts.

Since the DESI survey is based on optical ground-based data, it has a spatial resolution of roughly $1''$. For its counterpart association we start by investigating the entire DESI EDR catalogue (DESI Collaboration et al. 2024b), to which we apply the following set of selection criteria to obtain a subsample of objects with robust spectroscopic redshifts:

- objtype = TGT;
- deltachi2 > 10;
- zcat_primary = 1;
- coadd_fiberstatus = 0;
- zwarn < 4.

Furthermore, only spectra with good model fits and no serious issues with the redshift determination are used.

Out of the three EDFs, only the EDF-N overlaps with the DESI EDR. To obtain the counterparts, we set the maximum error to $1''$ based on DESI’s spatial resolution, and obtain a total number of 64 039 matches with the raw MER catalogues. After applying the quality cuts specified in Sect. 2.1.1, this number is reduced to 24 922 matches between DESI EDR and the quality-filtered EDF-N catalogue.

All the matched DESI spectra belong to the SV3 sample of the DESI EDR, which was covered by a larger number of passes than the yet to be released main survey of DESI. This means that these regions have a higher completeness and at times, due to stacking, deeper observations than what we can expect from the DESI main survey.

2.3.2. SDSS DR17

The Sloan Digital Sky Survey (SDSS, York et al. 2000) is a large-scale imaging and multi-fibre spectroscopic redshift survey that has mapped millions of objects from our Galaxy to the distant Universe, including stars, galaxies, and quasars. SDSS’ Data Release 17 (DR17) marks its fifth and final release from the fourth phase (Abdurro’uf et al. 2022). DR17 contains the entire release of the Mapping Nearby Galaxies at APO Survey (MaNGA, Bundy et al. 2015), as well as the MaNGA Stellar Library, and the complete release of the Apache Point Observatory Galactic Evolution Experiment 2 survey (APOGEE, Majewski et al. 2017). Moreover, DR17 also includes data from the SPectroscopic IDentification of ERosita Survey subsurvey (SPIDERS, Clerc et al. 2016; Dwelly et al. 2017) and the eBOSS-RM programme (Shen et al. 2015), as well as 25 new or updated value-added catalogues, covering a total of $14\,555\text{ deg}^2$ with an average PSF FWHM, typically measured in the r band, of around $1''.3$, and an approximate magnitude limit of around $r = 22.7$ (Blanton et al. 2017; Abdurro’uf et al. 2022). DR17 includes approximately 1.5 million unique spectra for sources, with available spectroscopic redshifts. The SDSS DR17 overlaps solely with the EDF-N, and even then, only a limited number of sources fall within this area. Nevertheless, we conduct a cross-match with the quality-filtered EDF-N catalogue using a fixed radius of $0''.5$, based on SDSS’ PSF FWHM, and obtain a total of 18 counterparts.

2.3.3. Other spectroscopic surveys

Gaia DR3 announced a sample of 6.6 million quasar candidates (the qso_candidates table¹; Gaia Collaboration: Valenari et al. 2023; Gaia Collaboration: Bailer-Jones et al. 2023), which has high completeness thanks to the combination of several different classification modules, including the Discrete Source Classifier (DSC), the Quasar Classifier (QSOC), the variability classification module, the surface brightness profile module, and the *Gaia* DR3 Celestial Reference Frame source table. Nevertheless, the *Gaia* DR3 QSO candidate catalogue has an estimated low purity of quasars (52%) and a large scatter of redshift estimates.

Instead of using the original *Gaia* DR3 QSO candidates catalogue, we take a purified version to find *Euclid* counterparts of the sources. This purified catalogue includes: (i) Quasia (Storey-Fisher et al. 2024), with nearly 1.3 million sources at $G < 20.5$; (ii) CatNorth (Fu et al. 2024), with more than 1.5 million sources

¹ The *Gaia* DR3 quasar candidate catalogue is available at the *Gaia* archive <https://gea.esac.esa.int/archive> with table name gaiadr3.qso_candidates.

down to the *Gaia* limiting magnitude in the 3π sky of the Pan-STARRS1 (PS1; Chambers et al. 2016) footprint ($\delta > -30^\circ$); and (iii) CatSouth (Fu et al., in prep), with 0.9 million sources with $G < 21.0$ covered by the fourth data release (DR4) of the SkyMapper Southern Survey (SMSS; $\delta \lesssim 16^\circ$; Onken et al. 2024). The compilation of the three catalogues contains more than 1.9 million unique (with unique *Gaia* source_id) quasar candidates in the entire sky. This catalogue from now on will be referred to as the ‘purified’ GDR3 QSO candidate sample (GDR3-QSOs). We cross-match the *Euclid* fields with this integrated GDR3-QSOs catalogue using the Q1 provided *Gaia* IDs and find 647, 811, and 513 matches in quality-filtered EDF-N, EDF-S, and EDF-F, respectively.

3. Identified populations

In this section, we present the two main populations identified in the *Euclid* data for this work: stars and AGN. While the primary objective is to investigate various diagnostics for compiling a comprehensive AGN candidate catalogue, stars significantly contribute to the contamination of AGN selection techniques. Therefore, developing an effective selection method to identify the stellar population within the Q1 data is crucial.

To identify AGN, we use traditional colour selection techniques and investigate new colour diagnostics using the limited labelled data obtained after cross-matching the *Euclid* catalogues. Additionally, we refer to other Q1 papers that also explore AGN detection techniques. All these approaches include using a combination of *Euclid*’s photometry, photometric and spectroscopic information from the matched data sets, spectral energy distributions (SEDs) fitting, morphological analysis, and machine-learning techniques.

However, despite the cross-matching between the various data sets and *Euclid*, we lack a substantial number of reliably labelled sources, excluding those from DESI. Consequently, we are unable to accurately quantify the purity and completeness of some of these methods. Follow-up work is necessary to further test the methodologies described here and to accurately quantify these parameters.

For the AGN diagnostics based on *Euclid*’s photometry, we use the template-fit fluxes provided by the Q1 catalogues. These fluxes are colour corrected, following the prescription outlined in Euclid Collaboration: Romelli et al. (2025).

Additionally, we examined the impact of correcting the fluxes for Galactic extinction using $E(B - V)$ values from the Galactic dust map from Planck Collaboration (2014). As shown in Galametz et al. (2017), the colour of the source SED can lead to significant correction variations. In this work we provide extinction corrections for two extreme cases, blackbody temperatures of 100 000 K and 5700 K. As the Galactic latitudes of the EDFs were selected to be in regions of low Galactic extinction, we find that these corrections have a minimal effect on the template-fit fluxes (approximately 5% variation). Consequently, we decided not to implement these corrections.

3.1. Stellar candidates

The proper motion and parallax of an object track its apparent transverse movement over time, as well as its shift in position against a distant background when viewed from different angles. Most stars within our Galaxy show measurable proper motions and parallaxes due to their proximity to us. Vice versa, distant objects like quasars or galaxies have negligible proper motions

and parallaxes. Therefore, tracking these two parameters is crucial when attempting to identify stars.

Gaia DR3 provides parallaxes and proper motions for around 1.46 billion sources, with a limiting magnitude of about $G \approx 21$ and a bright limit of about $G \approx 3$ (Gaia Collaboration: Vallenari et al. 2023). By identifying the *Gaia* counterparts and utilising the information provided by *Gaia* DR3, we have the necessary data to identify as stellar candidates those sources with significant proper motion and parallaxes. However, since the detection limit of *Gaia* ($G < 21$) is not as deep as that of *Euclid* ($I_E \leq 26.2$), for sources beyond $G \geq 21$, an alternative method is required for detecting stellar candidates.

The Q1 data release also includes catalogues with object classifications that provide the probabilities of an object being a star, a galaxy, or a QSO based on the source’s photometry. This is obtained by performing a supervised machine-learning method called Probabilistic Random Forest (PRF, Reis et al. 2019). Specifics on the method used can be found in Euclid Collaboration: Tucci et al. (2025), hereafter referred to as TM25. To summarise, the classifiers estimate the probability of objects belonging to a particular class and set a threshold that must be surpassed for an object to be classified into one of the groups. The advised threshold for objects to be considered as stars differs between EDF-N, EDF-S, and EDF-F, with values of 0.58, 0.68, and 0.68, respectively. Upon investigating the data, we decided that a threshold of 0.7 for all three fields provides a purer sample of stellar candidates

By combining information from both *Gaia* and the Q1 PRF, we are able to construct a comprehensive approach for identifying stellar candidates. However, to refine the star selection and ensure that no extended objects are recorded to have large proper motions and/or parallaxes, or are misclassified by the random forest, we impose an additional condition on the morphology to select only star-like objects. Therefore, only point-like sources (i.e., MUMAX_MINUS_MAG < -2.6) are considered stellar candidates (Euclid Collaboration: Romelli et al. 2025).

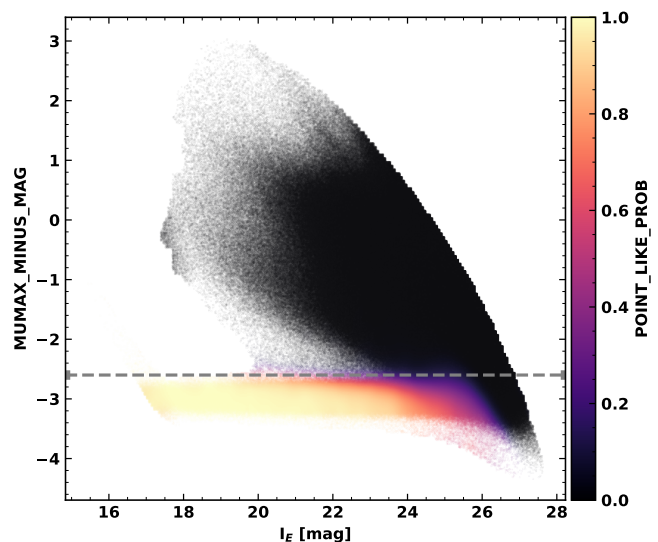


Fig. 2. MUMAX_MINUS_MAG versus I_E for sources in the EDF-N. The colour scale indicates the point like probability of a source. The dotted line indicates the threshold (MUMAX_MINUS_MAG < -2.6) below which most sources appear to be point-like.

The parameter MUMAX_MINUS_MAG is the difference between two quantities, both available in the Q1 catalogue: a global

Table 3. Number of selected stellar candidates per Q1 fields.

Field	No. stellar candidates		
	$18 < I_E \leq 21$	$21 < I_E \leq 22$	$22 < I_E \leq 24.5$
EDF-N	45 700	43 918	84 156
EDF-S	30 856	27 173	63 589
EDF-F	10 159	7 818	17 163

Notes. All sources must satisfy the conditions set by Eq. (1). The numbers reported are based on the quality-filtered catalogues.

measure, named MAG_STARGAL_SEP; and a local one, named MU_MAX. The first is the magnitude used to compute point-like probability, and the second is the peak surface brightness above the background detection level. MU_MAX measures the brightest, most concentrated light within a specific area of the source. The MUMAX_MINUS_MAG is useful to identify point-sources or nearly point-like sources (Euclid Collaboration: Romelli et al. 2025). Figure 2 indicates how this cut in the data set (MUMAX_MINUS_MAG < -2.6) is able to capture those sources that have a high probability of being a point-like source.

As a result, we adopt the following prescription to select stellar candidates,

$$\text{MUMAX_MINUS_MAG} \leq -2.6 \wedge$$

$$\begin{cases} \sqrt{\left(\frac{\varpi}{\sigma_\varpi}\right)^2} > 5 \vee \sqrt{\left(\frac{\mu_{\alpha^*}}{\sigma_{\mu_{\alpha^*}}}\right)^2 + \left(\frac{\mu_\delta}{\sigma_{\mu_\delta}}\right)^2} > 5 & \text{for } G < 21, \\ \text{phz_star_prob} > 0.7 & \text{for } G \geq 21, \end{cases} \quad (1)$$

where ϖ stands for the parallax of an object, σ_ϖ the error of this measurement, and $(\mu_{\alpha^*}, \mu_\delta)$ are the proper motion measured in the right ascension and declination positions, with their corresponding errors, $(\sigma_{\mu_{\alpha^*}}, \sigma_{\mu_\delta})$. We refer to these selected sources as `stellar_candidates`. Table 3 gives a summary of the number of stellar candidates that we identify in each of the three Q1 fields.

3.2. AGN candidates: photometric selection

We now present different AGN selections applied to the Q1 fields. This initial section examines previously established photometric criteria and introduces novel photometric diagnostics developed for this work. We report the number of AGN candidates recorded per criterion in Table 4.

3.2.1. Probabilistic random forest

As outlined in Sect. 3.1, Q1 provides object classification catalogues with the probabilities of an object being a star, galaxy, or QSO (TM25). They also provide the probability threshold to use in order to select different populations of sources. In particular, the recommended thresholds for QSOs are 0.67, 0.85, and 0.85 for the EDF-N, EDF-S, and EDF-F, respectively. When investigating the data we decided that a threshold of 0.85 for the EDF-N, and 0.95 for the EDF-S and EDF-F created an overall more refined sample of QSOs. However, since the PRF is trained using photometry only, to avoid contaminants from stellar objects that might have QSO-like colours, we exclude the stars identified by Sect. 3.1. We refer to this ‘purified’ version of the TM25 catalogue as the PRF candidates.

The total number of identified QSO candidates with this recipe in the quality-filtered catalogues is 180 666, and the breakdown per field is reported in Table 4.

3.2.2. Bisigello+24 selections

In preparation for the *Euclid* mission, Euclid Collaboration: Bisigello et al. (2024), from now on referred to as BL24, identified various colour-colour selection criteria for AGN using *Euclid* photometry alone, and combinations of *Euclid* photometry with additional external photometric bands. Their study was carried out for both, the EWS and the EDS, using simulated data from the Spectro-Photometric Realisations of IR-Selected Targets at all- z (Spritz, Bisigello et al. 2021). Selection criteria were identified by maximising the F1-score – the harmonic mean of the completeness (C , the fraction of true AGN recovered from the original AGN sample) and the purity (P , the fraction of true AGN among all AGN candidates).

In this work, due to the depth of the MER catalogues, we use the diagnostics derived for the EWS. It is worth noting that all selection criteria developed in BL24 assumed that stars had been previously selected and removed from the sample. Therefore, before applying any of their selections, we remove all stellar candidates identified in Sect. 3.1.

We then apply their purest selection criterion in the EWS to identify QSO candidates based on three *Euclid* filters. This selection provided them with an F1 score of 0.224 ± 0.001 , and was derived from a low purity ($P = 0.166 \pm 0.015$), as well as a low completeness ($C = 0.347 \pm 0.004$). This particular selection follows the prescription

$$(I_E - Y_E < 0.5) \wedge (I_E - J_E < 0.7) \wedge [I_E - J_E < -2.1(I_E - Y_E) + 0.9]. \quad (2)$$

This criterion identifies a total of 521 252 QSO candidates in the EDF-N, 755 032 in the EDF-S, and 315 683 in the EDF-F.

In addition to stars, based on the large number of candidates obtained and the low purity of this selection, it is possible that other contaminants also affect this selection. Therefore, to improve the purity of this diagnostic, which was specifically designed to identify unobscured AGN, also referred to as Type I AGN, we impose an additional requirement of point-likeness to eliminate potential extended contaminants (i.e., MUMAX_MINUS_MAG ≤ -2.6). This criterion removes approximately 87% of candidates in each field, indicating that many of the initial candidates were extended sources. However, by applying this additional morphological filter, we exclude several AGN at $z \leq 1.2$ that may appear as extended sources due to *Euclid*’s resolution.

Although it is challenging to assess how this additional condition might enhance the purity of the selection, this approach is a straightforward method to further clean the candidate sample without deviating from the prescription established by BL24. The combination of selecting sources with Eq. (2) and MUMAX_MINUS_MAG ≤ -2.6 will henceforth be referred to as selection ‘B24A’. The total number of QSO candidates obtained in the quality-filtered catalogues after applying the B24A selection is 211 797, the breakdown of which is reported in Table 4. Furthermore, Fig. 3 illustrates the QSO candidates selected with B24A in the EDF-N, showcasing the colours of all quality-filtered compact sources in this colour-colour plane and highlighting the QSOs selected as candidates.

BL24 also included some additional diagnostics designed to combine the *Euclid* photometry with that of other surveys. Specifically focussing on future surveys, such as Rubin/LSST, they characterised a series of criteria using the *ugriz* bands. Since the *Euclid* photometry is accompanied by ancillary ground-based optical photometry taken with these bands (details in

Table 4. Number of selected AGN candidates per criteria.

Data	Diagnostic	EDF-N			EDF-S			EDF-F		
		$18 < I_E \leq 21$	$21 < I_E \leq 22$	$22 < I_E \leq 24.5$	$18 < I_E \leq 21$	$21 < I_E \leq 22$	$22 < I_E \leq 24.5$	$18 < I_E \leq 21$	$21 < I_E \leq 22$	$22 < I_E \leq 24.5$
Photometry	PRF ^a	138	1 882	44 429	161	2 838	37 869	59	1 476	17 694
	B24A	1 085	2 038	12 146	1 210	2 113	13 155	480	886	4 690
	B24B	658	997	8 178
	C75	1 106	2 081	10 888	2 149	3 379	17 031	1 278	1 639	7 561
	R90	498	196	405	792	347	690	399	204	579
	GDR3-QSOs	606	35	3	769	42	0	442	68	0
	$JH_{-I_E}Y$	1 084	1 789	15 599	1 244	2 369	18 236	509	980	6 888
	I_EH_{-gz}	591	909	6 802	787	1 247	9 340	400	611	4 323
	Total	2 547	6 474	80 184	3 410	8 505	80 390	1 697	3 995	34 854
	Spectroscopy	QSO	400	484	673
Broad-line QSO		366	460	603
Broad-line Galaxy		57	48	113
NII BPT		93	1	1
SII BPT		69	0	5
OI BPT		72	3	38
WHAN		1 792	12	14
Blue		104	47	103
KEX		188	110	467
Total		2 250	636	1 350
Other QI selections	RW25 ^b	62	72	210	625	425	609	447	520	1 973
	SG25 ^c	1 831	2 078	1 326	3 435	2 678	1 609	1 400	1 012	571
	MB25 ^d	494	2 813	12 779	650	3 361	15 043	232	1 464	6 721
Total	2 303	4 566	13 700	4 179	5 691	16 169	1 759	2 471	7 325	...
SED fitting	LB25 ^e	2 172	1 218	4 166

Notes. ^(a) PRF selected QSOs: Euclid Collaboration: Tucci et al. (2025). ^(b) X-ray selected AGN candidates: Euclid Collaboration: Roster et al. (2025). ^(c) AGN candidates from diffusion models trained on VIS data: Euclid Collaboration: Stevens et al. (2025). ^(d) AGN candidates from deep learning trained on IllustrisTNG: Euclid Collaboration: Margalef-Bentabol et al. (2025). ^(e) AGN candidates with spectroscopic redshift from SED fitting: Euclid Collaboration: Laloux et al., in prep

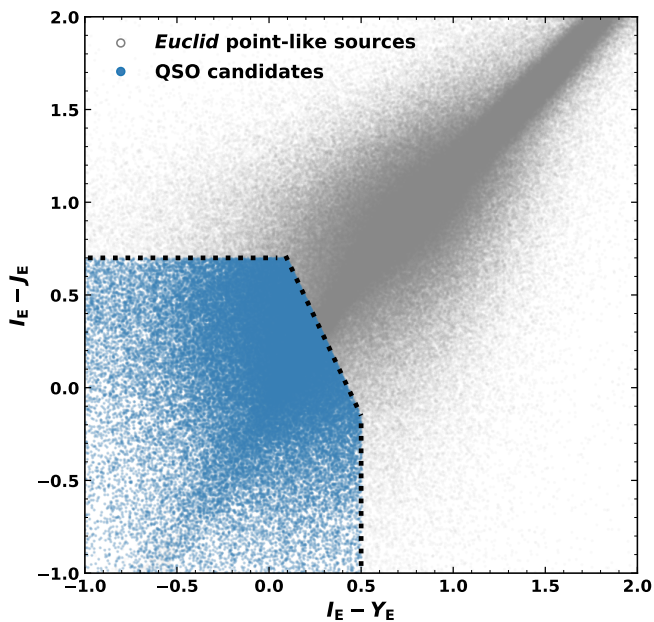


Fig. 3. B24A selection criteria applied to the EDF-N. In grey we show all quality-filtered *Euclid* compact sources, while the QSO candidates are shown in blue. We also showcase, with the black dotted line, the limit of the B24A selection.

RE25), we opt to test the selection criteria derived using the u and z bands.

This specific selection follows the prescription

$$(I_E - H_E < 1.1) \wedge (u - z < 1.2) \wedge [I_E - H_E < -1.3(u - z) + 1.9], \quad (3)$$

and was designed to identify Type I AGN, assuming all stellar candidates had already been removed from the samples. For this selection, BL24 obtained $F1 = 0.861 \pm 0.004$, with $C = 0.813 \pm 0.011$ and $P = 0.922 \pm 0.017$, making it the purest and most complete selection of their work.

The EDF-N is the only Q1 field that contains the u band as part of the *Euclid* ancillary data. Therefore we can only apply this selection to the EDF-N, obtaining a total of 1 092 763 QSO candidates.

Similarly to B24A, we suspect that the large number of candidates might be attributed to potential contaminants infiltrating this selection. Consequently, driven by these other contaminants, we apply the same morphological cut to this criterion. We refer to the combination of Eq. (3) and $\text{MUMAX_MINUS_MAG} \leq -2.6$ as the B24B selection. Applying this combination reduces the number of selected QSO candidates to 114 145. We report this number, split into magnitude bins, in Table 4. Additionally, Fig. 4 illustrates the QSO candidates selected with B24B in the EDF-N, showcasing the colours of all quality-filtered point-like sources and highlighting the QSOs selected as candidates. Once again, it is not possible to assess how the additional condition on morphology might impact the purity of the selection, but this approach provides a way of further cleaning the candidate sample without having to alter the BL24 criterion.

3.2.3. WISE-AllWISE selection

Utilising the supplementary photometry derived from the WISE-AllWISE counterparts, we implement the selection criteria established by Assef et al. (2018), hereinafter called A18. In their

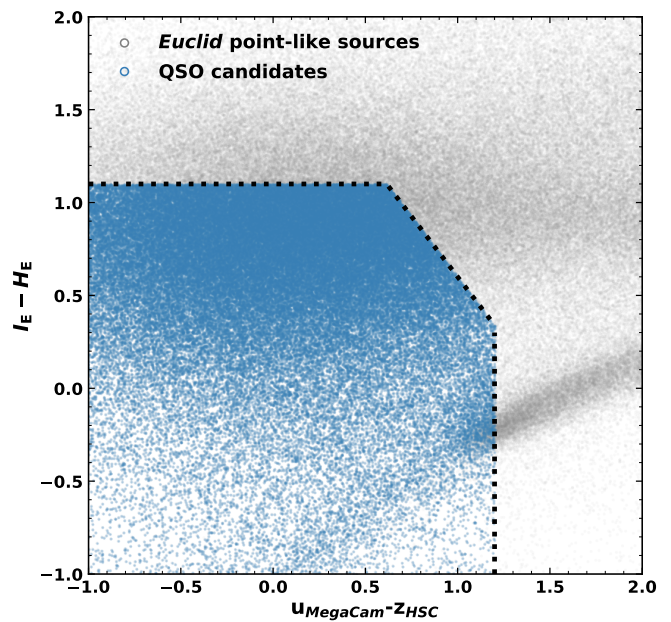


Fig. 4. B24B selection criteria applied to the EDF-N. In grey we show all quality-filtered *Euclid* point-like sources, while the QSO candidates are shown in blue. We also showcase, with the black dotted line, the limit of the B24B selection.

work, two distinct diagnostics, C75 and R90, were examined. As their names suggest, these criteria were designed to generate catalogues with 75% completeness and 90% reliability, whereby reliability measures the purity of the selection. The completeness-optimised AGN diagnostic is defined by

$$W1 - W2 > 0.71, \quad (4)$$

where the completeness fractions for a given $W1 - W2$ colour cut are independent of magnitude. Meanwhile, the reliability driven AGN selection takes the form,

$$W1 - W2 > \begin{cases} 0.65 \exp[0.153(W2 - 13.86)^2] & , \text{ for } W2 > 13.86, \\ 0.65 & , \text{ for } W2 \leq 13.86. \end{cases} \quad (5)$$

Additionally, to maintain the completeness and reliability of these diagnostics, it is essential to impose extra conditions, such as $W1 > 8$, $W2 > 9$, $\text{SNR}_{W2} > 5$, and the WISE-AllWISE quality flags $\text{cc_flags} = 0$. Both of these selections are established for the Vega magnitude system.

Before applying either one of these selections, we remove the stellar candidates identified with the Sect. 3.1 prescription. The total numbers of C75 and R90 AGN candidates is 65 083 and 4 688, respectively, and the numbers of AGN candidates per field are reported in Table 4. Figure 5 shows both of these selections applied to EDF-N sources, where we include all *Euclid* sources matched to WISE-AllWISE and we highlight those sources that are selected as AGN candidates by C75 or R90.

3.2.4. Gaia DR3

Using the *Gaia* DR3 cross-match data, we identify sources marked as QSO candidates from the GR3-QSOs sample. Specifically, we identify 647, 811, and 513 candidates in the quality-filtered catalogues of EDF-N, EDF-S, and EDF-F, respectively.

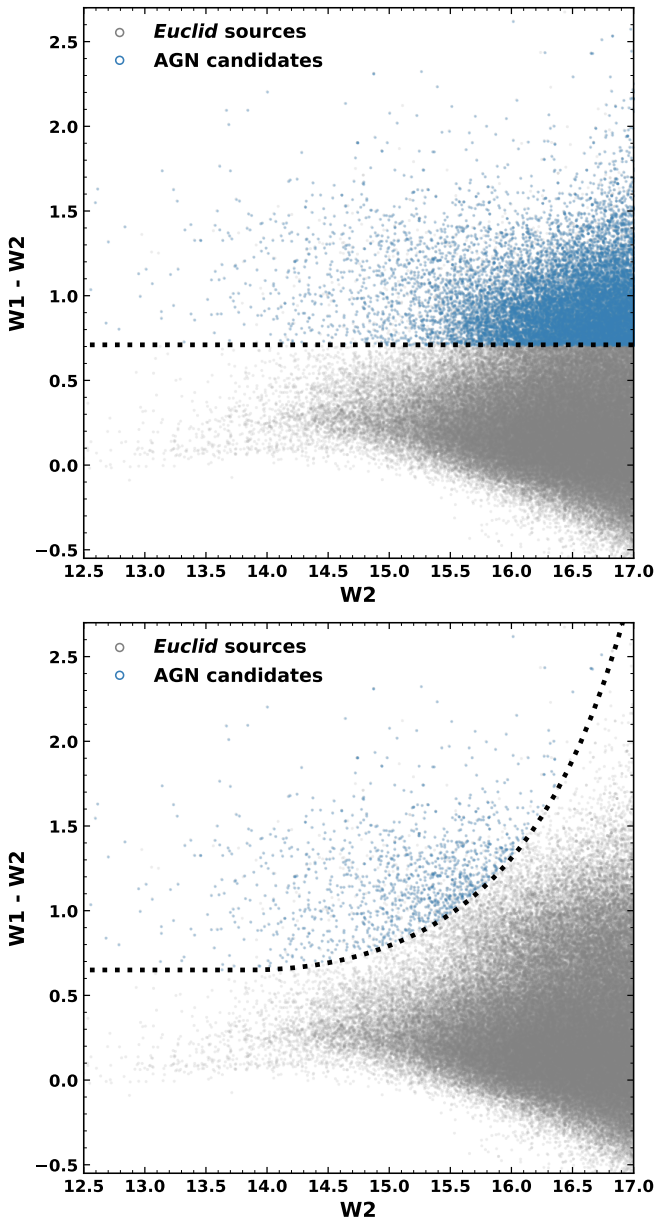


Fig. 5. WISE-AllWISE AGN candidates in the EDF-N defined by Eq. (4), top panel, and Eq. (5), bottom panel. In grey we show all *Euclid* sources with WISE-AllWISE counterparts, while the blue sources represent the selected AGN candidates. The black dotted lines represent the limits of the C75 and R90 selections. The stellar candidates have been removed from the samples of AGN candidates.

To address potential stellar contamination within this sample, we exclude the very few sources classified as stellar candidates according to Sect. 3.1. The refined number of sources in the quality-filtered catalogues is presented in Table 4.

3.2.5. New *Euclid*-only colour cut: $JH_{I_E}Y$

Euclid provides us with a myriad pieces of information that can be utilised for selecting QSO candidates. Motivated to obtain a purer QSO selection, we investigate a new diagnostic using all four *Euclid* bands, starting with a morphological cut. Focusing on QSOs, we consider only point-like sources with $MUMAX_MINUS_MAG \leq -2.6$. Imposing this restriction allows us

to exclude all extended sources that could act as contaminants in our selection, to the expense of detectable AGN within extended host galaxies (see Sects. 3.2.3, 3.3 and 3.4).

Subsequently, we create a smaller ‘ground truth’ catalogue where we include sources that have been matched to DESI and have been classified as DESI QSOs while possessing broad-line detection and z_{spec} (see Sect. 3.3.1). In this way, we are able to differentiate between objects that have a good reliability of being QSOs, galaxies, or stellar candidates.

Imposing good photometry on all sources (i.e., working with the quality-filtered catalogues), we explore the colour–colour space $J_E - H_E$ versus $I_E - Y_E$. We identify a cut that excludes the stellar locus, and obtain the following:

$$\begin{aligned}
 &[-0.1 \leq (I_E - Y_E) < 1.0] \wedge [-0.5 \leq (J_E - H_E) < 0.6] \\
 &\quad \wedge \{[(J_E - H_E) > 0.5(I_E - Y_E) - 0.20] \\
 &\quad \vee [(J_E - H_E) > 0.13]\}.
 \end{aligned} \tag{6}$$

We then investigate the confusion matrix, completeness and purity in two z_{spec} bins separated by $z_{\text{spec}} = 1.6$. We limit this analysis to $I_E < 21$, since at fainter magnitudes galaxy contamination from DESI (especially at $z > 1.6$) is largely unknown. The confusion matrix is computed applying the following labels to the ‘ground truth’ sample:

- true; DESI QSO labelled objects with a detected broad-line and spec-z are broad-line QSOs (BLQSOs); or
- false; anything that is not a broad-line QSO acts as contaminants (galaxies, Type II AGN, stars).

Then, depending on whether or not an object is within our colour-colour selection we can assign:

- positive; is compact and is within the selection; or
- negative; is not point-like or is outside the selection.

We opt to count stellar contaminants across all redshift bins, given that stars can interfere with selection processes at both low and high z . Using the aforementioned prescription, we achieve $P = 0.92$ with $C = 0.63$ for the $I_E < 21 \wedge z_{\text{spec}} < 1.6$ bin, and $P = 0.95$ with $C = 0.90$ for the $I_E < 21 \wedge z_{\text{spec}} > 1.6$ bin. Nevertheless, these values should be taken with caution since they can not be straightforwardly extrapolated to fainter magnitudes, as the presence of numerous contaminants, especially compact galaxies, could significantly reduce the purity and completeness of this selection at fainter magnitudes. Given the lack of sufficiently reliable galaxy labels at fainter magnitudes, assessing the impact of contaminants is challenging. Therefore, we consider this selection method particularly effective for our two brightest I_E bins, while the statistics for the faintest bin remain less constrained. In Appendix A, we illustrate the number of candidates picked up by this selection, split into our three magnitude bins, where it is evident that at fainter magnitudes, the number of candidates is considerably larger, hinting at higher contamination rates.

With this specific criterion, we identify a total of 313 714 QSO candidates, the breakdown of which can be found in Table 4. Additionally, Fig. 6 illustrates the application of this colour cut to the EDF-N, highlighting the corresponding point-like sources that served as DESI counterparts and were used to derive this QSO selection criterion, as well as the stellar locus that we identify in Sect. 3.1.

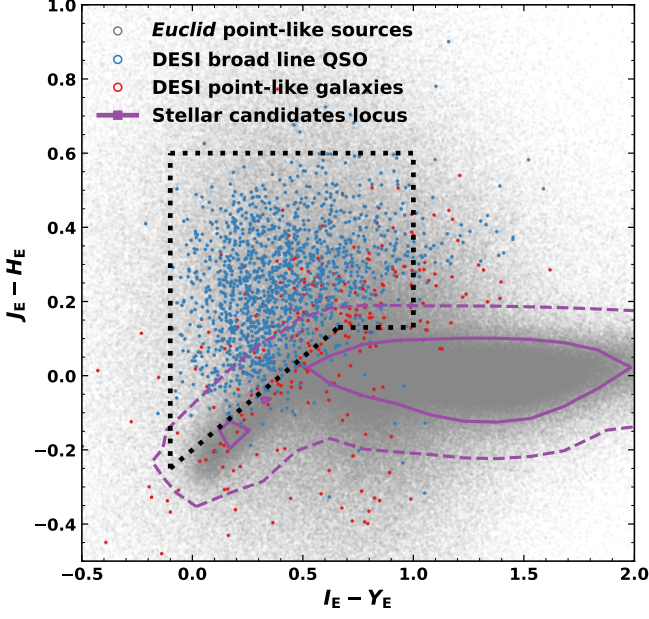


Fig. 6. New colour-cut criteria defined by Eq. (6), black dotted line, applied to EDF-N. In grey we show all *Euclid* compact sources that have a DESI counterpart. The blue coloured points represent the *Euclid* compact sources selected as DESI BLQSO candidates and the red ones represent the *Euclid* compact objects selected as galaxy candidates by DESI. Moreover, the purple lines represent the 68% (solid) and 95% (dashed) contours of the stellar candidates found in Sect. 3.1.

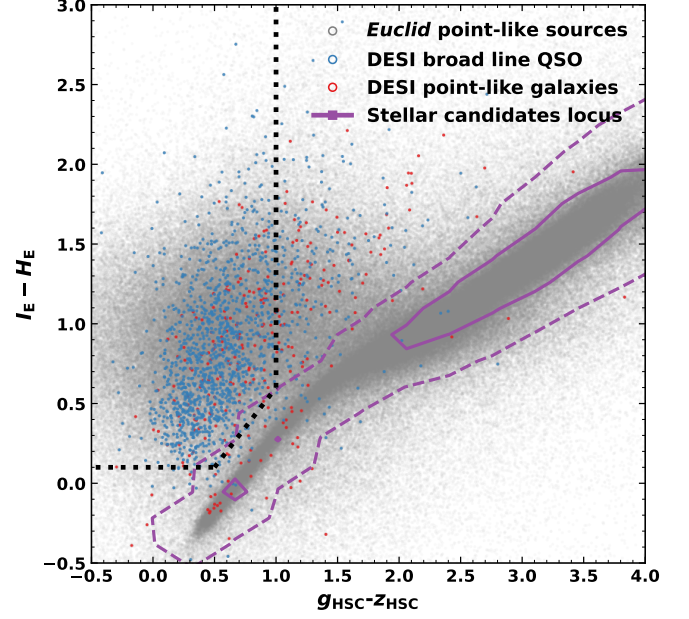


Fig. 7. New colour-cut criteria defined by Eq. (7), black dotted line, applied to EDF-N. In grey we show all *Euclid* compact sources with that contained a DESI counterpart. The blue coloured points represent the *Euclid* compact sources selected as DESI BLQSO candidates and the red ones represent the *Euclid* compact objects selected as galaxy candidates by DESI. Moreover, the purple lines represent the 68% (solid) and 95% (dashed) contours of the stellar candidates found in Sect. 3.1.

3.2.6. New *Euclid* and ancillary photometry colour cut: $I_E H_{gz}$

Employing *Euclid*'s ancillary data from Q1, we explore an additional diagnostic using the $I_E - H_E$ vs. $g - z$ colour space. Similarly to Sect. 3.2.5, we create a smaller 'ground truth' catalogue including sources that have been matched to DESI classified as DESI broad-line QSOs. We impose the same morphology cut as in Sect. 3.2.5 and we then identify the area occupied by the stellar locus, as well as the locations that the DESI broad-line QSOs and galaxies populate. Based on this initial realisation, it becomes apparent that it is possible to separate these objects in this colour space with the following prescription:

$$(g - z < 1) \wedge [(g - z < 0.5) \wedge (I_E - H_E > 0.1)] \vee [(g - z > 0.5) \wedge (I_E - H_E > (g - z) - 0.4)]. \quad (7)$$

We then follow the same guidelines as the ones presented in Sect. 3.2.5 and investigate the confusion matrix, C and P , in two z_{spec} bins separated by $z_{\text{spec}} = 1.6$, again limiting the analysis to $I_E < 21$. We obtain $P = 0.93$ with $C = 0.60$ for $I_E < 21 \wedge z_{\text{spec}} < 1.6$, and a $P = 0.97$ with $C = 0.77$ for $I_E < 21 \wedge z_{\text{spec}} > 1.6$. Once again, given the limited number of labels, particularly at fainter magnitudes, we refrain to assess the performance of this selection at $I_E > 21$. Appendix A illustrates this selection split into the three I_E bins to show the increasing number of candidates and therefore contaminants with fainter magnitudes.

With this specific criterion, we identify a total of 267 513 QSO candidates, its breakdown shown in Table 4. Figure 7 shows the application of this colour cut to the EDF-N, highlighting the corresponding point-like sources that serve as DESI counterparts and are used to derive this QSO selection criteria.

3.3. AGN candidates: spectroscopic selection

We present the various AGN selections applied to the Q1 fields based on spectroscopy obtained from the DESI counterparts. We report the number of AGN candidates recorded per criterion on Table 4.

3.3.1. DESI selection

We investigated the presence of QSOs and AGN within the 64 039 matches of DESI EDR to the *Euclid* matches. We began by creating a subsample of extragalactic objects that only included targeted objects with positive redshifts that had not been spectroscopically selected as DESI stars. To do so, we implemented the following criteria:

- $z > 0.001$;
- $\text{spectype} \neq \text{STAR}$.

From this subsample, the simplest method for selecting the QSO candidates is by using the DESI spectral type classification (SPECTYPE=QSO, DESI Collaboration et al. 2024b). Additionally, for sources classified as galaxies (SPECTYPE=GALAXY, DESI Collaboration et al. 2024b), we make use of multiple AGN diagnostics based on emission line fluxes, widths, and equivalent widths measured with FastSpecFit (Moustakas et al. 2023). These measurements are available for 40 274 of the DESI EDR-*Euclid* MER sources.

- The detection of broad $H\alpha$, $H\beta$, $Mg\ II$ or $C\ IV$ emission lines with a $\text{FWHM} \geq 1200 \text{ km s}^{-1}$.
- An AGN classification in either the $N\ II$ ($[O\ III]\lambda 5007/H\beta$ versus $[N\ II]\lambda 6583/H\alpha$), $S\ II$ ($[O\ III]\lambda 5007/H\beta$ versus $[S\ II]\lambda 6717, 6731/H\alpha$) or $O\ I$ ($[O\ III]\lambda 5007/H\beta$ versus

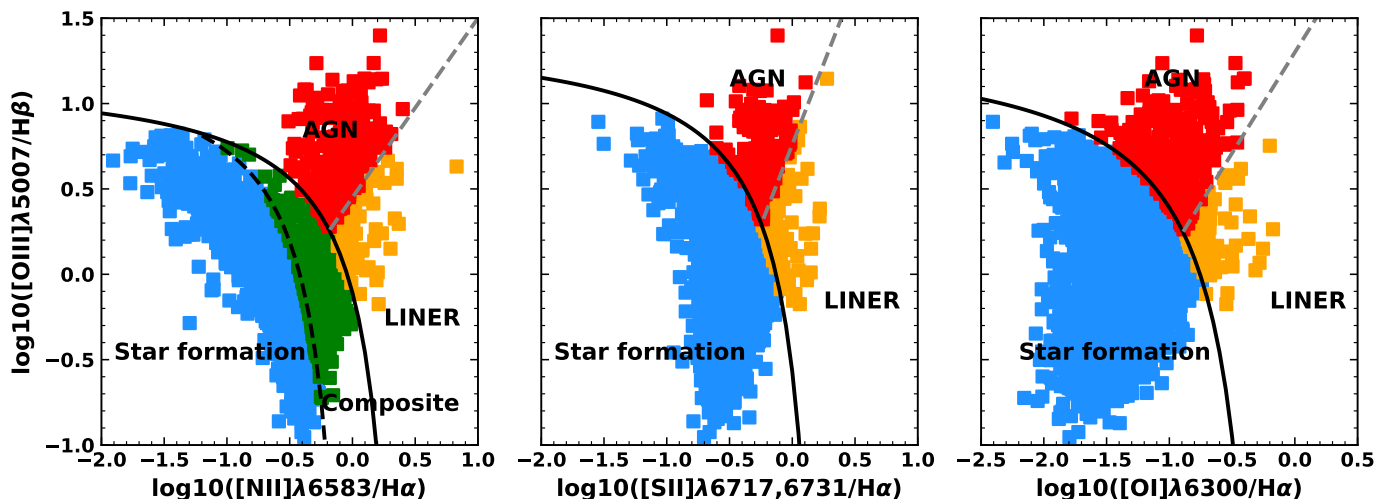


Fig. 8. N II (left), S II (centre), and O I (right) emission line diagnostic diagram for those *Euclid* sources with a DESI spectroscopic counterpart.

[O I] λ 6300) emission line diagnostic diagrams (or ‘BPT diagrams’ after Baldwin et al. 1981, see Fig. 8). For the N II BPT, we make use of Kewley et al. (2001), Kauffmann et al. (2003), and Schawinski et al. (2007) to distinguish between AGN, Low-ionization nuclear emission-line region (LINER), composite, and star formation (SF) ionisation. For the S II and O I BPTs, we make use of the Kewley et al. (2006) and Law et al. (2021) demarcations to distinguish between AGN, LINER and star-formation ionisation.

- A strong AGN or weak AGN classification in the WHAN diagram of Cid Fernandes et al. (2010), which makes use of the equivalent width of the H α emission line.
- An AGN classification in the BLUE diagram of Lamareille (2010), which makes use of the equivalent width of the H β and [O II] λ 3727 emission lines.
- An AGN classification in the KEX diagnostic diagram of Zhang & Hao (2018), which makes use of the [O III] λ 5007 emission line width.

Moreover, to identify QSOs containing broad emission lines, we also perform a test to detect broad H α , H β , Mg II, or C IV emission lines with a FWHM ≥ 1200 km s $^{-1}$ for those sources classed with SPECTYPE=QSO.

This results in a total of 4392 AGN candidates in the quality-filtered catalogue, the breakdown of which is shown in Table 4. Additionally, Fig. 8 shows an example of the N II, S II, and O I emission-line diagnostics performed for the DESI counterparts.

3.4. AGN candidates: other AGN selections

The following AGN diagnostics were developed in other works conducted in preparation for the Q1 data release. We report the number of AGN candidates recorded per criterion on Table 4.

3.4.1. X-rays

Euclid Collaboration: Roster et al. (2025), subsequently referred to as RW25, present the Q1 counterparts to X-ray point sources, starting from the 4XMM-DR14 (Webb et al. 2020), the *Chandra* Source Catalog (CSC) Release 2 Series (Weisskopf et al. 2002; Evans et al. 2024) and the eROSITA (Predehl et al. 2021) first Data Release (DR1; Merloni et al. 2024). For each of the Q1 fields and each of the X-ray surveys they first identify the best

Euclid counterpart by means of the Bayesian algorithm NWAY (Salvato et al. 2018), which assigns the probability of a good association considering a) the separations between sources, their positional uncertainties, and their number density, and b) the similarity between the SED of a candidate counterpart and the SED of a typical X-ray emitter, regardless whether the source is Galactic or extragalactic. The latter information is provided by a prior externally defined using a random forest on a set of *Euclid*-only features (i.e., no features from ground-based photometry) extracted from the Q1 catalogues for a training sample of secure X-ray emitters and secure field sources. The same procedure is repeated, randomizing the coordinates of the X-ray catalogues so that the probability of a chance association can be determined (see RW25 for details).

After the determination of the counterparts, the authors then assign to each source a probability of being Galactic (star, compact object) or extragalactic (galaxy, AGN, QSO). This is again done using a training sample of secure Galactic and extragalactic sources and *Euclid*-only features from the Q1 catalogues.

Finally, for the sources that have a probability larger than 50% of being extragalactic, the authors provide either photometric redshifts computed using PICZL (Roster et al. 2024) on Legacy Survey DR10 (Dey et al. 2019) images (and thus limited to the sources detected in that survey), or spectroscopic redshifts from literature.

The released catalogues enable users to refine their samples based on specific scientific needs, balancing purity and completeness through the NWAY output parameters. In total, they report 12 645 AGN candidates, though some sources have multiple counterparts; when considering only the best match for each unique X-ray source, the sample reduces to 11 286 candidates. They identify 949 in EDF-N, 3789 in EDF-S, and 6548 in EDF-F.

From their catalogue, in order to identify those candidates that have the highest probability of being an AGN candidate, we select a subsample of sources with $P_{\text{Gal}} < 0.2$. Implementing our ‘cleaning’ on these candidates we obtain 434, 1812, and 3813 X-ray candidates in the quality-filtered EDF-N, EDF-S, EDF-F catalogues (see Table 4).

3.4.2. Diffusion models

Euclid Collaboration: Stevens et al. (2025), from here on called **SG25**, use the reconstruction error of a diffusion model, a type of generative model, to select AGN candidates. Through training on VIS images, the model is able to learn a bias for the light profile at the centres of galaxies. Since AGN are rare, the bright pixel and steep fall off of light they exhibit is converted to one that is significantly flatter, leading to a high reconstruction error for suspected AGN. They obtain a total of 15 940 AGN candidates across the three EDFs (see [Table 4](#)).

3.4.3. Deep learning

Euclid Collaboration: Margalef-Bentabol et al. (2025), hereafter called **MB25**, present a deep learning (DL) method to quantify the AGN contribution (f_{PSF}) of a galaxy using VIS imaging. The DL model is trained with a sample of mock images generated from the IllustrisTNG simulations, designed to mimic *Euclid* VIS observations, with different levels of AGN contributions artificially injected as PSFs. The DL model is trained to estimate the level of the injected PSF, achieving a root mean square error (RMSE) of 0.052 on the test set. After applying this model to the Q1 data, they find 48 840 galaxies across the EDFs that are classified as AGN based on this AGN contribution, that is, $f_{\text{PSF}} > 0.2$. Adopting a less conservative threshold of $f_{\text{PSF}} > 0.1$ increases the number to 158 711 AGN. This method allows the identification of AGN even when the AGN component is not the primary contributor to the host galaxy's luminosity. The resulting number of AGN candidates found per field and magnitude bin is in [Table 4](#).

3.5. AGN fraction from SED fitting

SED-template fitting is a powerful method for measuring physical properties of galaxies and AGN by reproducing the observed photometry with a combination of theoretical and empirical SED models that account for the different AGN and galaxy emission processes. The result of the multi-component SED fitting constrains a variety of physical properties, notably the AGN fraction, defined as the ratio of the AGN flux to the total flux in the MIR band, used for AGN identification (see e.g., [Dale et al. 2014](#); [Thorne et al. 2022](#)). By decomposing the emission from the galaxy and the potential AGN components, SED fitting permits the identification of fainter AGN compared to colour-colour approaches. Moreover, since the IIR emission is not significantly impacted by AGN obscuration, SED fitting can reliably identify obscured AGN missed by optical or X-ray methods ([Pouliasis et al. 2020](#); [Andonie et al. 2022](#)).

Since accurate redshift measurements are required for reliable SED fitting, we first restricted the analysis to sources with spectroscopic redshift in EDF-N, computing the AGN fraction as the ratio between the AGN flux and the total flux over the 5–20 μm wavelength range (following [Dale et al. 2014](#); [Thorne et al. 2022](#)). These results can be used to define an accurate selection threshold for AGN that could later be applied to the entire sample with photometric redshift. We used the SED fitting algorithm CIGALE ([Boquien et al. 2019](#); [Yang et al. 2022](#)) to fit the UV-to-mid-IR photometry of our sources, including: GALEX FUV/NUV, *ugriz*, *Gaia*-G/BP/RP, *Euclid* $I_E/Y_E/J_E/H_E$, WISE W1/2/3/4. Our modelling consists of a delayed star-formation history with a simple stellar population from [Bruzual & Charlot \(2003\)](#) and the [Chabrier \(2003\)](#) initial mass function, a galactic dust attenuation ([Calzetti et al. 2000](#)) and emission

([Draine et al. 2014](#)), nebular lines ([Inoue 2011](#)), and an AGN model ([Fritz et al. 2006](#)). For further details, we refer to *Euclid* Collaboration: Laloux et al., in prep, from now on referred to as LB25, which presents the physical properties of the AGN candidates in the three EDFs.

The results are shown in [Fig. 9](#), where the normalised cumulative distribution of the AGN fraction for normal galaxies is compared to the different AGN samples. As indicated by the vertical black dotted line, we define our AGN fraction threshold as the intersection between the normal galaxy distribution (dash-dotted blue line) and the broad-line AGN one (BLAGN, dash-dotted red line), whereby BLAGN refers to those sources classified as QSOs or galaxies that exhibit broad emission lines of H α , H β , Mg II, or C IV. An AGN fraction threshold at $f_{\text{AGN}} = 0.25$ is a compromise between purity and completeness, since it correctly selects 77% of the spectroscopically confirmed BLAGN while only 23% of the non-AGN candidates are misclassified as AGN. With this threshold we obtain a total of 7766 AGN candidates within the quality-filtered catalogues, the breakdown of which can be found in [Table 4](#). A lower AGN fraction threshold, such as $f_{\text{AGN}} = 0.1$ ([Thorne et al. 2022](#)), would lead to an increased AGN completeness (97%), but with a much higher false-detection probability (78%). Conversely, a higher AGN fraction threshold of $f_{\text{AGN}} = 0.5$ improves the purity, with a false-detection probability of 5%, at the expense of the completeness (44%).

In [Fig. 9](#), we also compare the AGN fraction distribution for different AGN-selection methods. We find that AGN candidates selected by C75, B24B, $I_E H_{-gz}$, and $JH_{-I_E Y}$ show a similar AGN fraction distribution to the BLAGN, while the R90 candidates tend to have higher AGN fractions. This suggests that the more reliable methods select sources where the AGN emission dominates over its host, potentially missing weaker AGN. Conversely, the B24A method tends to select sources with lower AGN fractions, some of which are likely to be contaminants, as indicated by the curve of the non-AGN candidates. Additionally, the cyan curve, representing the narrow-line AGN (NLAGN) from the DESI emission-line diagnostics, is significantly shifted to lower AGN fraction values compared to other methods. This demonstrates that while SED fitting is efficient at identifying unobscured AGN, achieving both high completeness and purity for obscured AGN is more challenging. Nevertheless, it still manages to reach 32% completeness for these elusive AGN.

In this Section we discussed only qualitatively the comparison of the results of the SED fitting analysis for different selection approaches applied to sub-samples with spectroscopic redshift. Not necessarily the same results will hold for samples which have only photometric redshifts as the availability of a spectroscopic redshift could impact the selections and introduce a bias in the AGN fraction distributions. The AGN fraction from SED fitting for the three deep fields for spectroscopic and photometric redshift sources will be provided in LB25.

4. Discussion

4.1. Contaminants

Although *Euclid* offers exciting prospects for AGN selection, similar to all AGN criteria, several contaminants must be considered when developing these methods. As already discussed in [Sect. 3.1](#), stars are a big contaminant when it comes to colour-colour diagnostics. However, other types of stellar objects also make it hard to create a pure AGN-selection technique.

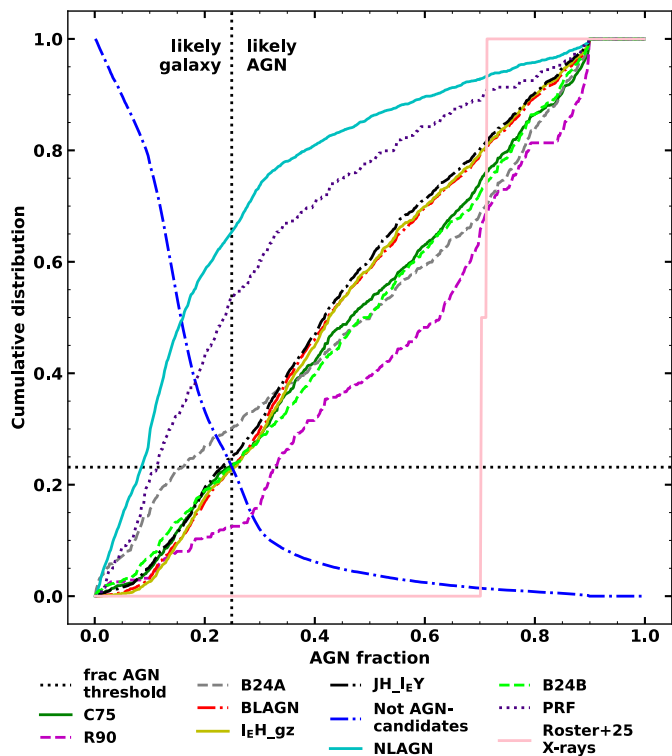


Fig. 9. Normalised cumulative distribution of the AGN fraction measured by SED fitting for the different AGN selection methods in EDFN: C75/R90 (green solid/purple dashed), B24A/B (grey/lime dashed), spectroscopically-confirmed broad/narrow line AGN (red dash-dotted/cyan solid), I_eH_{gz} (yellow solid), JH_{I_eY} (black dash-dotted), PRF (dark blue dotted) and X-ray (pink solid). The blue dashed-dotted line is the inverse cumulative distribution of all non-AGN candidates. The vertical black solid line represents the proposed AGN fraction threshold ($f_{AGN} = 0.25$) discriminating AGN and non-AGN sources, while the horizontal black dotted line shows the corresponding false-positive probability (23%). Note that only two X-ray sources have counterparts with z_{spec} in the quality-filtered EDF-N catalogue.

For instance, young stellar objects (YSOs), asymptotic giant branch (AGB) stars, and H II-regions can show very similar colours to those of AGN, and therefore slip through the selection criteria as potential candidates.

Another form of contaminants, related to low mass products of star formation would be that of brown dwarfs (Davy Kirkpatrick et al. 2011) whose feature at $1 \mu\text{m}$ can sometimes resemble a Ly α break at redshifts $z = 6-7.5$ (Wilkins et al. 2014; Hainline et al. 2024; Langeroodi & Hjorth 2023).

Compact normal galaxies can also play a big role as contaminants in AGN diagnostics (Kouzuma & Yamaoka 2010). This is because, at higher redshifts, distant galaxies may appear as point-like sources and show similar colours to that of AGN, therefore contaminating the AGN locus in colour-colour plots.

Additionally, dwarf irregulars (Irr) have a 4000 \AA break, meaning that, at some redshifts ($z < 1$), their SEDs are very similar to those of QSOs (BL24).

Moreover, high- z star-forming galaxies (SFGs) can sometimes also be considered a contaminant, since their SEDs resemble those of AGN. However, since their 4000 \AA break lies within *Euclid*'s bands at $z > 1$, it might be easier to separate between AGN and SFGs at these higher redshifts.

These contaminants are likely affecting the colour selections analysed in this work. With our current knowledge and under-

standing of the Q1 data, we are unable to identify these sources, making it difficult to distinguish between them and potential AGN candidates. Consequently, we acknowledge that, especially at fainter magnitudes, these sources are definitely contaminating our AGN sample, which worsens our calculated values for purity and completeness. Addressing this issue would require further work, such as combining our data with other photometric catalogues that include bands that would facilitate the separation of these objects, or obtaining more spectroscopic data to test for specific emission lines, like those mentioned in Sect. 3.3.1.

4.2. Comparison to expectations

Prior to this work, Euclid Collaboration: Selwood et al. (2025), from now on addressed as SM25, examined the AGN surface density expected for the *Euclid* mission. Starting from an X-ray luminosity function (Fotopoulou et al. 2016), they predicted the observational expectations for AGN with $z < 7$ in the EWS and EDS. They generated volume-limited samples covering $0.01 \leq z \leq 7$ and $43 \leq \log_{10}(L_{bol}/\text{erg s}^{-1}) \leq 47$. Each AGN was assigned an SED based on its X-ray luminosity and redshift. Dust extinction was applied, and once assigned and scaled, they performed mock observations of each AGN SED in their sample, convolving with the *Euclid* bands and an assortment of ancillary photometric bands to explore the observable population of $z < 7$ AGN in the *Euclid* surveys. They concluded that *Euclid* should be able to detect significantly more AGN in the EWS and EDS compared to those identified in other surveys covering similar regions. In the EDS they predicted an unobscured AGN surface density of 346 deg^{-2} based on the BL24 Eq. (2) selection, which is comparable to the densities obtained by *Spitzer* (Lacy & Sajina 2020) and the XXL-3XLSS (Chiappetti et al. 2018) survey. Similarly, the EWS was estimated to recover an AGN surface density of 331 deg^{-2} , surpassing the AGN densities obtained from both ground-based optical and space-based MIR missions.

We compare the AGN surface densities recovered in $18 < I_e \leq 24.5$ and $18 < I_e \leq 22$ for each one of the selections investigated in this work against the EWS predictions from SM25, which were based on a 5σ limit in all four *Euclid* bands, corresponding to a limiting magnitude of $I_e \approx 26.2$.

Starting with the BL24 candidates, in Fig. 10 we observe differences in the surface densities recovered for B24A and B24B. Given that B24B could only be applied to the EDF-N, the area used to calculate its AGN surface density is that of the EDF-N (23.9 deg^2), as opposed to the Q1 area (63.1 deg^2), which is used for most of the other selections. B24A/B have distinct statistics, with B24A having $P \approx 0.17$ and $C \approx 0.35$, while B24B has $P \approx 0.92$ and $C \approx 0.81$. Based on these numbers, the AGN surface densities for both selections align with their expectations. B24A retrieves a greater number of AGNs, likely due to its low purity, indicating that a notable amount of contaminants could be identified as candidates. The notably lower AGN surface density obtained for B24B can be attributed to the high purity of this selection, which indicates a smaller number of contaminants. Both selections reveal a significant difference in AGN surface densities recovered in $18 < I_e \leq 24.5$ and $18 < I_e \leq 22$, highlighting that at fainter magnitudes, we encounter a higher number of contaminants that are difficult to distinguish from actual AGN.

From Fig. 10 it is apparent that the number of AGNs selected by the PRF is comparable to those selected by B24A in the range of $18 < I_e \leq 24.5$, but it is significantly smaller in the range of $18 < I_e \leq 22$. This variation is likely due to the PRF potentially misclassifying compact galaxies with QSOs at fainter magnitudes. The larger number of candidates at fainter

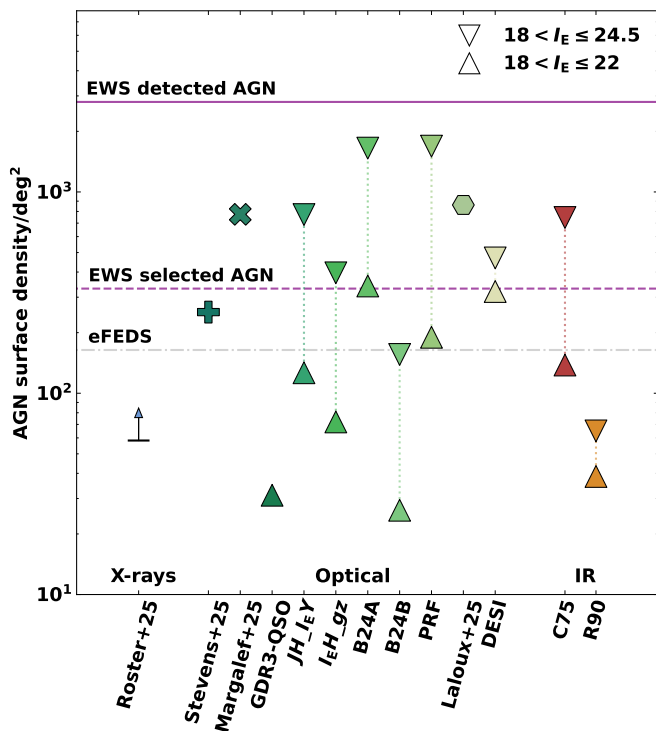


Fig. 10. Comparison of AGN surface densities obtained from the selection methods discussed in this work, divided into energy bands: X-ray selections (blue); optical selections (shades of green); and IR selections (shades of orange). For the selections B24A, B24B, $JH_{I_E}Y$, I_EH_{gz} , C75, R90, DESI, and PRF the AGN surface densities are split into $18 < I_E \leq 24.5$ (upside down triangles); and $18 < I_E \leq 22$ (upside up triangles). For the GDR3-QSOs, only the $18 < I_E \leq 22$ is shown due to the limiting magnitude of *Gaia*. Individual markers indicate AGN surface densities from other Q1 related works, including X-rays from RW25 (with the lower limit indicated by an arrow pointing up); morphology-based selections from SG25 (plus sign); and MB25 (cross); and SED fitting from LB25 (hexagon). The predictions for the detectable AGN (purple horizontal solid line) and identifiable AGN (purple dashed line) in the EWS from SM25 are included. The grey horizontal dashed line represents recovered AGN surface density by eFEDS (Liu et al. 2022a).

magnitudes indicates that, even when we increase the thresholds beyond the recommended values, many contaminants still affect this selection, highlighting that the advised Q1 PRF thresholds are too lenient and should be revised.

The observed numbers for both the C75 and R90 criteria align with what we expected for both of these catalogues. For the C75 criteria, with a completeness of 75%, we anticipated a larger number of candidates, as illustrated in Fig. 10. However, this expectation comes with the potential for an increase in contaminants. By excluding the stellar candidates, we have likely marginally increased the reliability of this selection, though we currently lack the tools to precisely quantify these effects. In contrast, R90 aims to create a more reliable, or purer, catalogue, resulting in a smaller AGN candidate size, as illustrated in Fig. 10. Notably, C75 recovers almost over an order of magnitude more candidates than R90 in both $18 < I_E \leq 24.5$ and $18 < I_E \leq 22$. Another observation from R90 is that the AGN surface density for $18 < I_E \leq 24.5$ and $18 < I_E \leq 22$ are not very different from each other, whereas other methods show a significantly higher higher AGN surface density in the faintest bin. The relatively bright limiting magnitude of this WISE-AllWISE selection ($W1 < 17.0$ (Vega) = 19.7 (AB)), together with its high pu-

riety, makes this selection highly incomplete at faint I_E magnitudes, since the number of very red selected candidates (with faint I_E magnitude) is small. This effect produces the small increase of the number of candidates going from $I_E < 22.0$ to $I_E < 24.5$.

Both of the new selections, $JH_{I_E}Y$ and I_EH_{gz} , recover similar AGN surface densities, which in the $18 < I_E \leq 24.5$ appear to surpass the expectations of what is identifiable as an AGN according to SM25. However, the high purity achieved by these selections was only quantified for $I_E < 21$. Therefore, we consider the AGN surface densities recovered in $18 < I_E \leq 22$ to be more reliable since we lack the means to assess how the P and C of these selections might change with increasing magnitude. This raises questions about the reliability of the candidates obtained for both selections at $I_E > 21-22$. Further work involving additional galaxy labels or spectra may be necessary to verify whether these AGN surface densities are accurate or inflated by contaminants.

The GDR3-QSO sample was originally quite small compared to the other samples of candidates (1971), so it was expected that the recovered AGN surface density was going to be orders of magnitude lower than that of other surveys. In Fig. 10, we only show GDR3-QSO's number density for $18 < I_E \leq 22$. This is due to the lack of GDR3-QSO candidates at the faintest magnitudes, which is linked to *Gaia*'s detection limit at $G < 21$. Additionally, DESI only covers a reduced area of the EDF-N, which we calculate to be approximately 9 deg^2 . We use this area to obtain the corresponding AGN surface density for the sources selected with DESI, which include QSOs, galaxies with detected broad-lines, and AGN selected via BPT or other narrow-line emission diagnostics. We find that the results are consistent with the expectations outlined by SM25 within the range $18 < I_E \leq 22$, and exceed the expectations for $18 < I_E \leq 24.5$. This is indicative of potential synergies between *Euclid* and DESI in the future when more data from both survey are available. Moreover, we also assess the overall AGN surface density of the X-ray AGN identified by RW25, which is lower than that recovered from eFEDS (Liu et al. 2022b). However, we note that the X-ray catalogue created in RW25 is a combination of different X-ray catalogues with varying depths (see Figure 1 of RW25). Therefore, the surface density reported in Fig. 10 should only be taken as a lower limit, where to estimate the surface area, we generated a multi-order coverage map with a resolution of $6''87$, yielding a total area of 36.72 deg^2 . We also note that the SED fitting AGN surface density from the LB25 candidates surpasses the limit of the identifiable AGN set by SM25. This could be due to the latter's use of Type I specific diagnostics to obtain their predictions, while SED fitting is more efficient, though not perfect, at identifying both, obscured and unobscured AGN. However, SED fitting is still affected by contaminants, and the quality of the available photometry will dictate the quality of the SED fit.

For most of the selections applied in this work, there seems to be a consistent trend. The magnitude range $18 < I_E \leq 24.5$ recovers a larger AGN surface density, likely filled with contaminants, while the brightest $18 < I_E \leq 22$ bin shows lower densities, usually below the predictions of SM25. Despite the goal of achieving a reliable surface density across all magnitudes, further work is required to assess and enhance the reliability of these selections in the faintest magnitudes. Therefore, we consider the $18 < I_E \leq 22$ range to provide a 'purer' catalogue of AGN candidates.

The resulting AGN surface density of our selections B24A, B24B, $JH_{I_E}Y$, I_EH_{gz} , GDR3-QSOs, PRF, SED-fitting, and

DESI candidates is 3641 deg^{-2} for $18 < I_E \leq 24.5$. By applying this magnitude cut, we eliminate saturated sources from the brightest magnitudes as well as the faintest sources in the Q1 catalogues. We acknowledge that this approach results in missing a population of *Euclid* sources detectable at a 5σ limit of I_E , for which we currently lack the tools to study comprehensively. Given the potential contamination even within this cut, we propose that the purest sample of AGN lies in the magnitude range $18 < I_E \leq 22$, resulting in an AGN surface density of 482 deg^{-2} . Even after narrowing the magnitude range of our selection, the AGN surface density recovered remains higher than what was expected to be identifiable with *Euclid*, yet falls short of the expected number of detectable AGN. This indicates that further refinement of our selection criteria is necessary to bring the number of reliable AGN candidates closer to what should be observable. Notably, this need for improvement is especially critical for Type II AGN, which have been largely excluded from most of the photometric selections evaluated in this study (not including the MIR, X-ray, SED-fitting, and spectroscopic diagnostics). With machine-learning and contributions such as those conducted by SG25 and MB25 we hope to be able to bridge the gap between detected and selected AGN, thereby reducing the bias against Type II AGN that typically arises from most colour-colour selections. In fact, the recovered AGN surface densities from SG25 and MB25, which are limited to $I_E < 22$ and $I_E < 24.5$ respectively, already highlight how effective machine learning can be in this regard.

4.3. Comparison among AGN selections

In this work, we have examined a variety of AGN diagnostics to construct the first *Euclid* multi-wavelength catalogue of AGN candidates. However, as expected, most of these AGN selections are incomplete and biased.

Focusing on the photometric selections applied to the *Euclid* photometry, which include B24A, B24B, $JH_{I_E Y}$, and $I_E H_{gz}$, it is the latter diagnostic that achieves the highest purity, with $P \approx 0.93$ for $z_{\text{spec}} < 1.6$ and $P \approx 0.97$ for $z_{\text{spec}} > 1.6$. However, $JH_{I_E Y}$ ($P \approx 0.92$ for $z_{\text{spec}} < 1.6$ and $P \approx 0.95$ for $z_{\text{spec}} > 1.6$) and B24B ($P \sim 0.92$) are only slightly lower. For the new selections, $JH_{I_E Y}$ and $I_E H_{gz}$, the calculation of P and C is limited to $I_E < 21$, thus considering only the brightest sources. This limitation is problematic at the faintest magnitudes, where we can not assess the P and C values and the effects of different contaminants due to the lack of reliable labels (see Figs. A.3 and A.4).

To determine whether the new selections align with other QSO and AGN candidates, we examine the agreement of these methods with the GDR3-QSO and X-ray AGN candidates. From Figs. 11 and 12, which illustrate the GDR3-QSO and X-ray candidates in the selections $JH_{I_E Y}$ and $I_E H_{gz}$ colour spaces, it is evident that both selections agree with the GDR3-QSOs, capturing most of them, while only selecting some of the X-ray sources. This outcome is expected as both of the new selections are designed to select Type I AGN, i.e., QSOs, hence the agreement with GDR3-QSOs, while X-ray samples contain a significant fraction of Type II AGN, which would occupy a different region of the colour space. The agreement with the GDR3-QSOs can also be attributed to the fact that these sources are selected using *Gaia* information, which has a detection limit of $G < 21$, therefore somewhat agreeing with the $I_E < 21$ cut we use to calculate the P and C values of selections A and B.

The issue of contaminants at fainter magnitudes also affects the B24A and B24B selections. At the faintest magnitudes, these

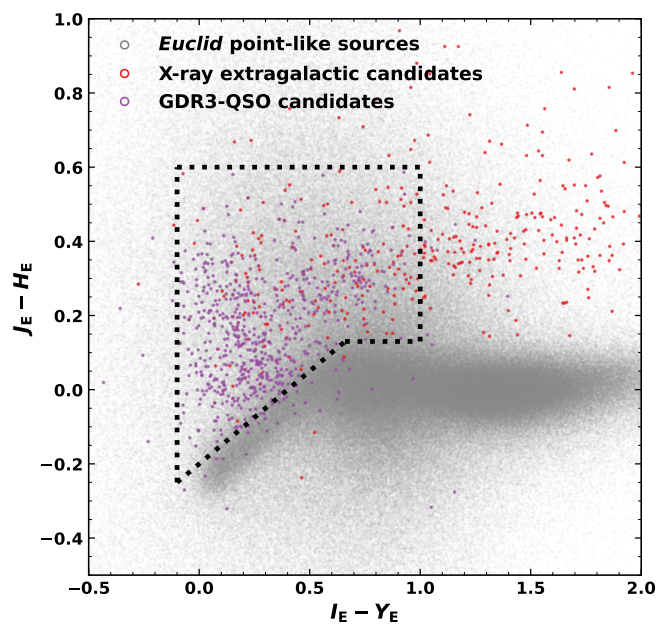


Fig. 11. Selection JH versus $I_E Y$ (black dotted line) applied to the EDF-N. In grey we show all *Euclid* compact sources. The red coloured dots represent the X-ray selected AGN candidates from RW25, while the purple dots indicate the GDR3-QSOs, which mainly lie within the selection.

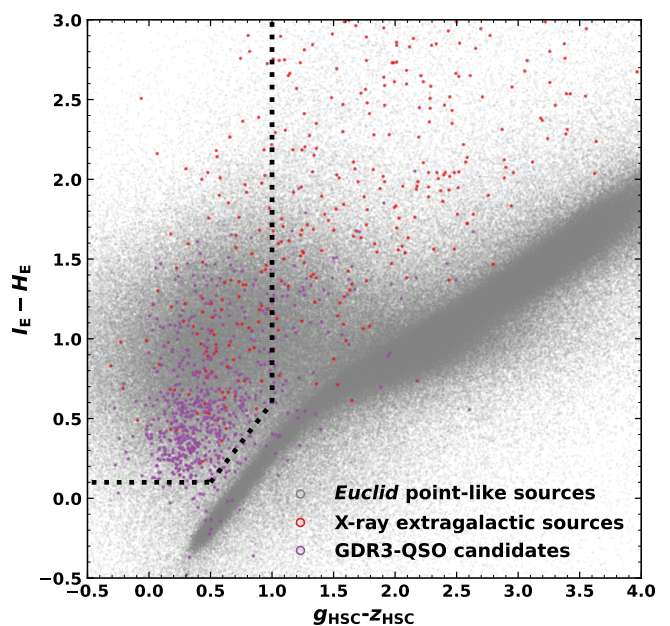


Fig. 12. Selection $I_E H$ versus gz (black dotted line) applied to the EDF-N. In grey we show all *Euclid* compact sources. The red coloured dots represent the X-ray selected AGN candidates from RW25, while the purple dots indicate the GDR3-QSOs, which mainly lie within the selection.

selections introduce a large number of candidates that are likely compact galaxies or other types of contaminants, such as brown dwarfs or stellar objects, which cannot be disentangled using colour cuts alone (see Figs. A.1 and A.2).

The DESI-selected QSO and AGN candidates exhibit higher reliability since they utilise DESI spectra from *Euclid* counterparts to assess the population a source may belong to. We partic-

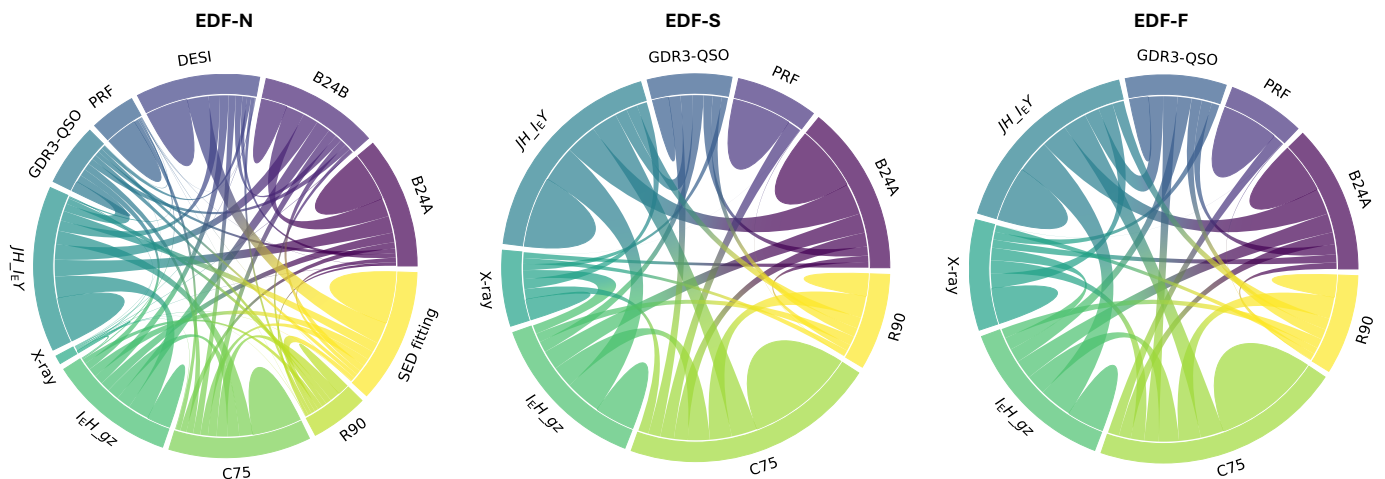


Fig. 13. Comparison among the number of sources selected as AGN candidates for the different criteria investigated in this work per EDF. We exclude the morphology-based AGN candidates due to the significant differences in their methodologies compared to other techniques explored in this work. We set the detection limit to the range of $18 < I_E \leq 22$, ensuring all selections match in depth. Note that EDF-N shows more selection criteria than EDF-S and EDF-F.

ularly trust those DESI BLQSO candidates, since these are indicative of Type I AGN activity, due to the high velocities of the ionised clouds within the broad-line region (BLR) of an AGN (Antonucci 1993; Veilleux 2002). Moreover, the NLAGN candidates tend to have high reliability because BPT diagnostics use a combination of nebular emission lines to differentiate between various ionisation mechanisms in gas (Baldwin et al. 1981). This helps distinguish between AGN, LINERs, SFGs, and composite objects, which encompass starburst-AGN objects (Kewley et al. 2006). The WHAN, KEX, and BLUE diagrams are similar in the sense that they also use specific line ratios and compare these to source qualities in order to identify different populations. However, although these may be more reliable AGN diagnostics at times, they are still incomplete and biased, and the sources identified with them should still be considered candidates.

For the specifics on the purity and completeness of the GDR3-QSOs, C75, R90, and X-ray selected AGN candidates, as well as those obtained using morphology and machine-learning information, we point the readers to the corresponding papers Storey-Fisher et al. (2024), Fu et al. (2024), Fu et al., in prep, A18, RW25, SG25, TM25, MB25, and LB25.

To assess the overlap among the various selection methods investigated (excluding the morphology-based ones, since their methodologies differ significantly from the other techniques explored in this work), Fig. 13 visualizes the portion of AGN candidates identified by multiple selections simultaneously per EDF. To enable a better comparison, we set the detection limit to the range of $18 < I_E \leq 22$, ensuring all selections match in depth.

We find that although most selections overlap to some extent in AGN candidates, there is a large number of candidates that do not co-exist in the different AGN samples. Notably, the PRF, B24A, and JH_IeY selections have substantial portions of their QSO candidate populations that are not selected by other diagnostics, potentially indicating that these sources are contaminants. Similarly, this behaviour is observed for some C75 candidates. However, the C75 selection is designed to identify AGN in general, which suggests that these sources may be Type II AGN, not detectable by the other QSO-specific diagnostics. Potential future work combining *Euclid*'s photometry with that of WISE-AllWISE could be a promising approach to reduce the

bias against Type II AGN. Nevertheless, this lies outside the scope of the current work.

We observe that most of the DESI AGN, GDR3-QSO, and X-ray extragalactic candidates appear to be consistently identified by the other diagnostics. Appendix B provides a numerical representation of Fig. 13 to quantify the agreements between selections.

In total, and including the AGN candidates identified in this work, which include B24A, B24B, JH_IeY , IeH_gz , DESI, PRF, SED-fitting and GDR3-QSOs, our current catalogue includes 229 779 AGN candidates in $18 < I_E \leq 24.5$, which is equivalent to an AGN surface density of $3\,641\text{ deg}^{-2}$. However, due to contamination, we believe the purest sample of AGN is in the magnitude range $18 < I_E \leq 22$, resulting in a total of 30 422 AGN candidates i.e. 482 deg^{-2} . This sample, although primarily composed of Type I AGN, also includes some Type II AGN identified through the DESI diagnostics.

4.3.1. Obscured vs. unobscured AGN

The majority of the AGN selection methods used in this work are specifically designed to select QSOs. These sources are easier to detect due to their distinct colours and point-like appearance. Additionally, as they are face-on AGN, they are minimally affected by dust, meaning their observed fluxes have not been significantly attenuated. The main drawback of optical photometric selections is that they are heavily affected by dust, making them easily optimised for Type I AGN, while being heavily biased against Type II AGN. Despite *Euclid* having NIR filters, it was predicted that selecting all AGN, including optically obscured AGN and composite systems, would be challenging with *Euclid* filters alone or supplemented by optical or other bands (BL24).

To ascertain this, we investigate the regions that the DESI spectroscopically selected narrow line AGN populate in different colour spaces. Figure 7 showcases an example of one of the DESI spectroscopic tests, the $N\text{ II}$ BPT diagnostic, plotted on the $I_E - H_E$ versus $g - z$ space. It is apparent that the area populated by the AGN is significantly entangled with that of composite, SFGs, and LINER galaxies, highlighting that obscured AGN do not occupy a specific and distinct region of the colour space. This

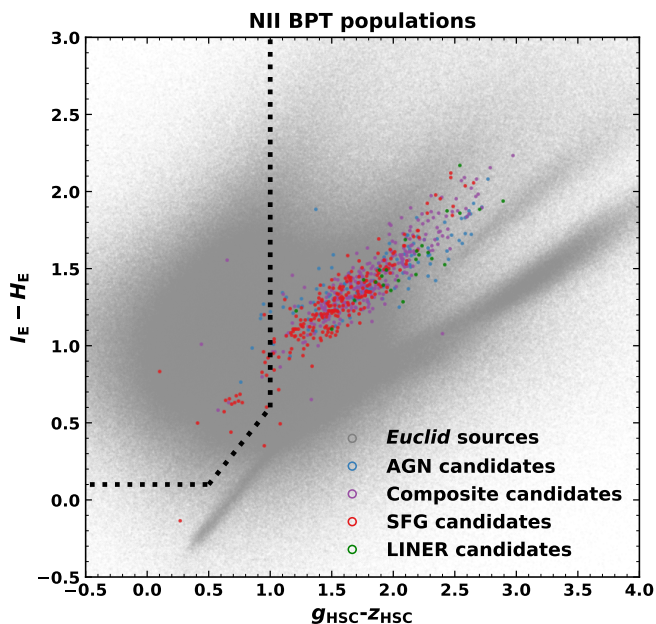


Fig. 14. Selection $I_E H$ versus gz (black dotted line) applied to the EDF-N. In grey we show all *Euclid* sources. The blue dots represent the N II selected AGN candidates, the composite galaxies are shown in purple, the SFGs in red and the LINERs in green.

behaviour was observed for all spectroscopic diagnostics across the different *Euclid* colour combinations.

Moreover we note that out of the total number of DESI broad-line QSOs (1 434) 91% are selected by our QSO diagnostics, while from the total of NLAGN selected emission line diagnostics (2 761) only 8% are detected by other selections, therefore highlighting the still existing bias against Type II AGN. Additionally, to further demonstrate the differences in populations identified by the spectroscopic and QSO selections, we use the ancillary photometry from the multi-wavelength catalogue to explore the SEDs of the AGN selected by these methods.

We first create a subsample of sources simultaneously identified as QSO candidates by B24A, B24B, $JH_{I_E Y}$, $I_E H_{gz}$, and DESI BLQSO, resulting in a total of 279 candidates. For these sources, we perform SED fitting with CIGALE using ancillary optical-to-IR photometric information and DESI's z_{spec} and examine their best-fit models and corresponding VIS cutouts. Figure 15 provides an example of the SED fitting and VIS images for two of these selected Type I QSO.

The SEDs show a notably strong AGN component in the mid-IR, even dominating the SF emission at longer wavelengths for source 2663093395657204902. Linked to the SF, we also notice moderate dusty absorption on the stellar emission. The corresponding AGN fractions and stellar masses obtained for these sources are approximately 0.46 and $M_* \approx 10^{10.5} M_\odot$ for source 2703664071644731101 and 0.69 with $M_* \approx 10^{10.1} M_\odot$ for source 2663093395657204902. The VIS cutouts reveal the point-like appearance of these candidates.

We then create a second subsample of sources identified as AGN candidates by either one of the DESI narrowline spectroscopic selections, and similarly fitted their optical-to-IR SED. Figure 16 showcases the SEDs and VIS cutouts for two of these AGN candidates. The corresponding AGN fractions and stellar masses obtained for these sources are approximately 0.31 and $M_* \approx 10^{10.5} M_\odot$ for source 2679391751656821227 and 0.97 with $M_* \approx 10^{10.8} M_\odot$ for source 2661957306668287025. It should

be noted that the absence of a WISE-AllWISE counterpart, and therefore photometry, for source 2679391751656821227 results in a large uncertainty (± 0.16) in its AGN fraction. The VIS cutouts reveal extended sources with bright centres and, in the case of 2661957306668287025, a dust lane.

This test highlights the difference between the average unobscured sources selected with our current QSO diagnostics and the obscured sources identified with spectroscopy. This opens an exciting path for future work to exploit similar types of spectroscopic diagnostics on the *Euclid* spectra to verify if more obscured AGN can be identified using *Euclid*'s spectroscopic capabilities, notably in extended sources, currently excluded of most QSO-selection approaches.

4.4. Galaxy major mergers with AGN contributions

Euclid Collaboration: La Marca et al. (2025) make use of a convolutional neural network (CNN) to perform morphological classification of a stellar mass-complete sample of Q1 galaxies in the redshift range $0.5 \leq z \leq 2.0$. The CNN is trained with *Euclid* I_E mock observations created from the IllustrisTNG simulations, with different levels of AGN contributions injected in 20% of the sample. The authors classify 113 155 galaxies as mergers and 269 933 as non-mergers. Then, they utilise the AGN catalogue presented here to select AGN in four different ways (X-ray detections, optical spectroscopy, through the f_{PSF} parameter, and with MIR colours) to study the possible connection of mergers with each AGN type. *Euclid* Collaboration: La Marca et al. (2025) observe a larger fraction of AGN in mergers compared to non-mergers, with the largest AGN excess seen in MIR-selected AGN, and a dependence of the merger fraction on the f_{PSF} parameter and the AGN luminosity. Their analysis supports the scenario in which mergers are most closely connected to the most powerful and dust-obscured AGN.

5. The AGN catalogue

The catalogue created through this work and made public through Zenodo contains the counterparts to the *Euclid* sources from GALEX, *Gaia*, WISE-AllWISE, DESI, SDSS, DES, and *Spitzer*, with their corresponding IDs, RA, and Dec (columns 1–25). We also include the flags to clean the data similarly to our work (column 26), split the data into our magnitude bin (27–29), and identify the stellar candidates (30–32). Additionally, columns 33–48 flag the sources that have been selected as AGN candidates via the various tests conducted in this study, including B24A, B24B, $JH_{I_E Y}$, $I_E H_{gz}$, C75, R90, PRF, QDR3-QSOs, and the different DESI diagnostics. Finally, we also include the results from the SED fitting explored alongside this work, which includes columns for the AGN fraction of those sources with DESI redshifts, their corresponding errors, and the resulting selected AGN candidates (columns 49–51). A detailed description of the columns included in this catalogue can be found in Appendix C.

6. Conclusions

In this paper, we have created and presented a multi-wavelength AGN candidate catalogue, incorporating ancillary photometric and spectroscopic data from surveys such as *Gaia*, WISE-AllWISE, DES, SDSS, DESI, and *Spitzer*. We summarise the most important results as follows.

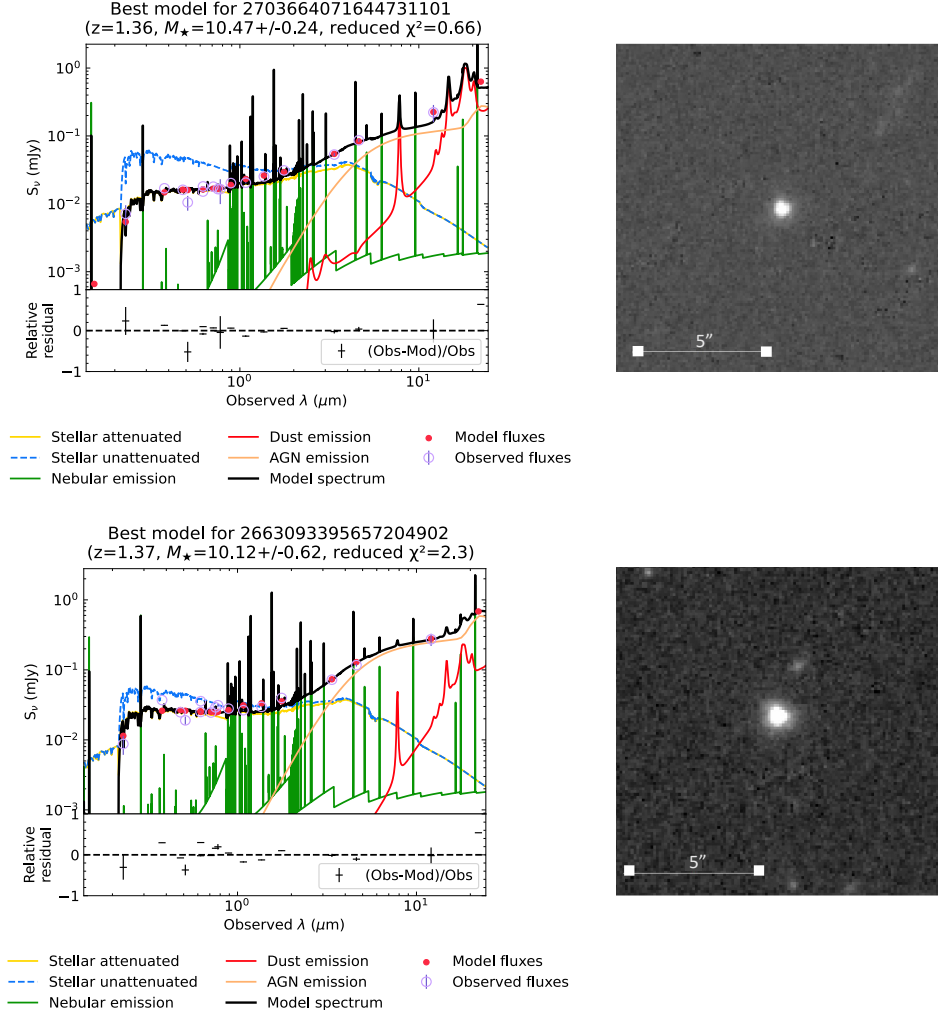


Fig. 15. SEDs and corresponding VIS cutouts of two QSO candidates identified by the selections JH_{I_EY} , I_EH_{gz} , B24A, B24B and simultaneously classified as broad-line QSOs by DESI.

- Counterpart associations are performed using a nearest-neighbour approach with STILTS, deciding the best fixed error radius for each survey based on their angular resolution and PSF FWHM.
- Two QSO diagnostics derived by BL24 are applied to the Q1 data. Upon investigating and refining these selections with a morphology cut ($MUMAX_MINUS_MAG < -2.6$), we obtain a total of 211 797 QSO candidates using *Euclid* photometry only (B24A), and 114 145 QSO candidates using *Euclid* plus ancillary bands (B24B).
- We apply the C75 and R90 A18 diagnostics to the WISE-AllWISE counterparts and obtain a total of 65 083 and 4 688 AGN candidates respectively.
- Labelled sources from the DESI counterparts are used to create two new QSO diagnostics: one tailored for *Euclid*-only photometry (JH_{I_EY}), achieving $P \approx .92$ and $C \approx .63$ for $z_{\text{spec}} < 1.6$ and $P \approx .95$ and $C \approx .9$ for $z_{\text{spec}} > 1.6$; and another using *Euclid* plus ancillary data (I_EH_{gz}), obtaining $P \approx .93$ and $C \approx .60$ for $z_{\text{spec}} < 1.6$ and $P \approx .97$ with $C \approx .77$ for $z_{\text{spec}} > 1.6$, both for $I_E < 21$.
- The spectra of DESI counterparts are utilised to test for broad-line detection in the DESI QSOs and galaxies. Additionally, spectroscopic tests (N_{II} BPT, S_{II} BPT, O_{I} BPT, WHAN, BLUE, and KEX) are performed to identify a sample of obscured narrow line AGN. In total we obtain 4392 DESI AGN candidates.
- Our catalogue is matched to a ‘purified’ version of the *Gaia* DR3 QSO catalogue (GDR3-QSO) with 1971 QSO candidates.
- By conducting SED fitting on a subset of sources with available z_{spec} , we are able to determine their AGN fraction and consequently identify a total of 7766 AGN candidates.
- Results are compared to other Q1 AGN-related works *Euclid* Collaboration: Tucci et al. (2025), *Euclid* Collaboration: Roster et al. (2025), *Euclid* Collaboration: Stevens et al. (2025), *Euclid* Collaboration: Margalef-Bentabol et al. (2025), and *Euclid* Collaboration: Laloux et al., in prep, assessing the differences and strengths of each selection.
- The purity and completeness of our selections are discussed, acknowledging the need for more labels in order to assess the impact of contaminants at the faintest magnitudes.
- A total of 229 779 AGN candidates are identified at $18 < I_E \leq 24.5$, with a refined sample of 30 422 AGN candidates within the magnitude bin of $18 < I_E \leq 22$.
- The AGN surface density expected from SM25 in the EWS, 331 deg^{-2} , is compared to our catalogue, $3 641 \text{ deg}^{-2}$, which reaches a higher AGN surface densities, most probably due to contaminants in the faintest magnitudes. Even when limiting AGN to $18 < I_E \leq 22$, with 482 deg^{-2} , we surpass the ex-

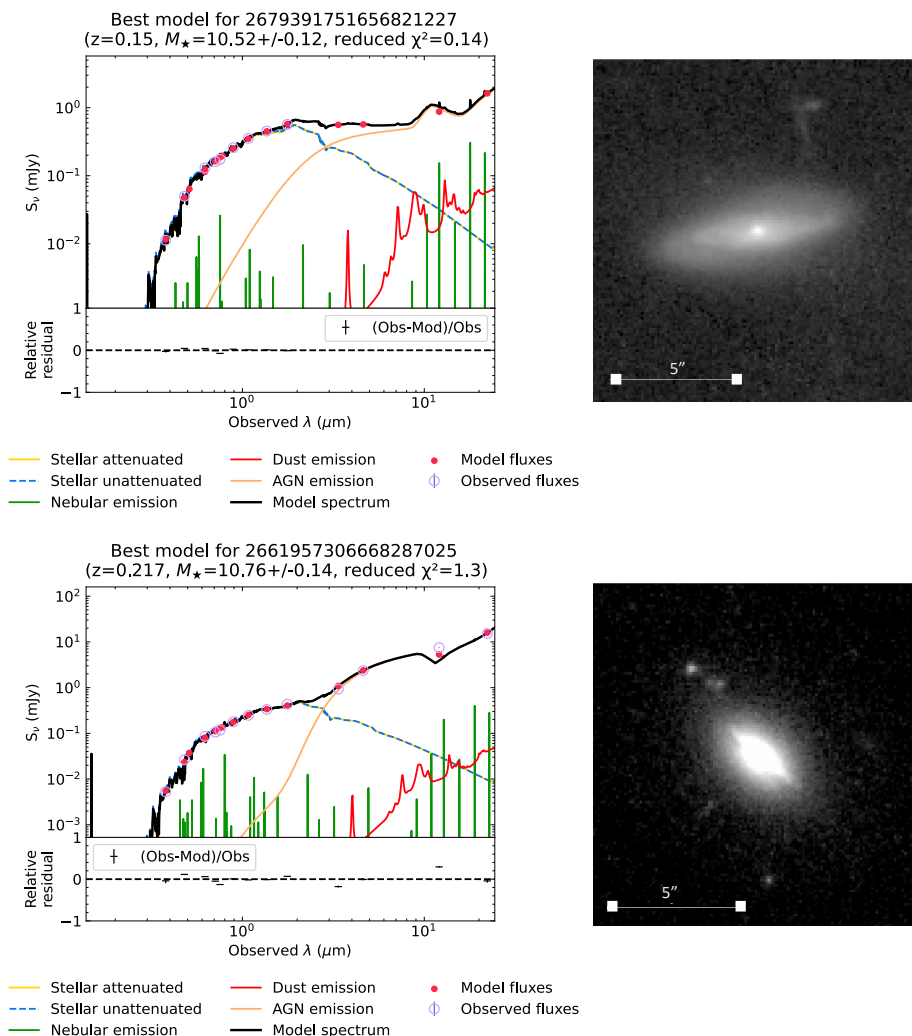


Fig. 16. SEDs and corresponding VIS cutouts of two narrow line AGN candidates identified via BPT diagnostics.

pected number of selected AGN, although we still fall short of the expected detected AGN. This gap could be bridged by future machine-learning studies.

- The AGN catalogue is presented, containing a wealth of information, including the data needed to replicate the numbers obtained in this work, as well as flags to easily identify different types of selected AGN.

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Based on data from UNIONS, a scientific collaboration using three Hawaii-based telescopes: CFHT, Pan-STARRS, and Subaru www.skysurvey.cc. Based on data from the Dark Energy Camera (DECam) on the Blanco 4-m Telescope at CTIO in Chile <https://www.darkenergysurvey.org>. This work uses results from the ESA mission *Gaia*, whose data are being processed by the *Gaia* Data

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Appendix A: QSO candidates in magnitude bins

To highlight the increase in contaminants with increasing magnitudes, we plot the selected QSO candidates across three magnitude bins for the *Euclid*-based photometric selections.

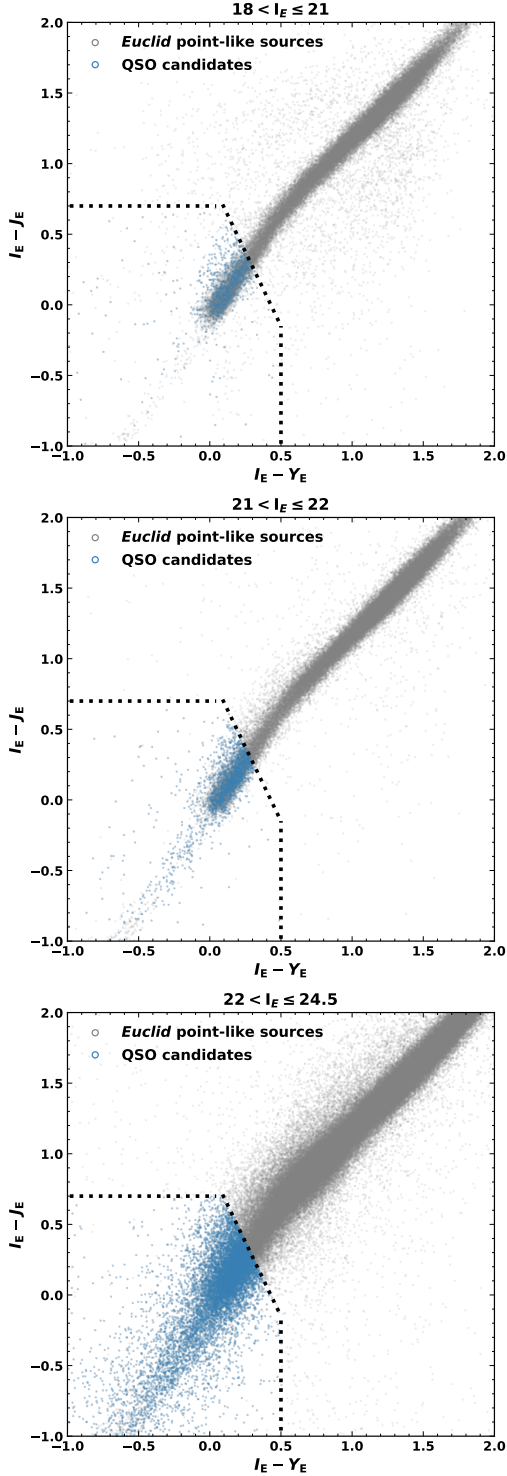


Fig. A.1. Comparison between the number of B24A QSO (blue) candidates per magnitude bin in the EDF-N. In grey we show all *Euclid* compact sources in the corresponding magnitude bin. Because the magnitude bins go from brighter colours (top plot) to faintest (bottom plot) the number of sources and QSO candidates increase.

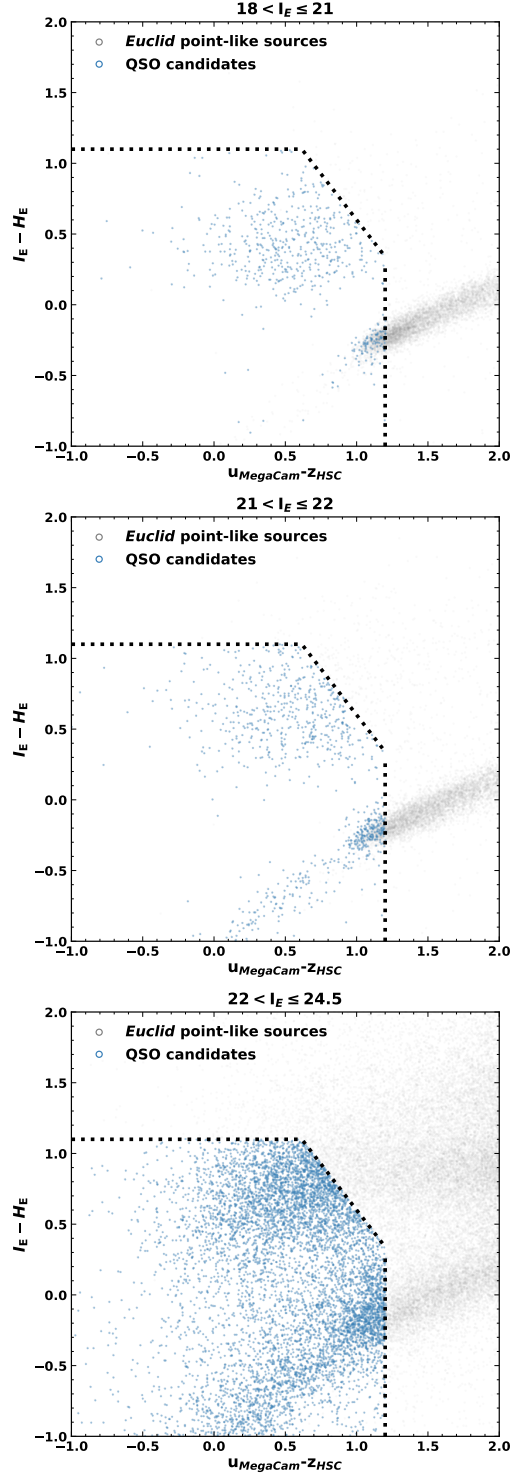


Fig. A.2. The same as Fig. A.1, but for B24B

Figure A.1 showcases selection B24A, Fig. A.2 B24B, Fig. A.3 $JH_{I_E}Y$, and Fig. A.4 I_EH_{gz} . It is evident that each selection is impacted by the higher number of candidates at fainter magnitudes. This highlights the necessity to conduct further work at $I_E > 21$, either by refining our selections or devising methods to identify potential contaminants in this region.

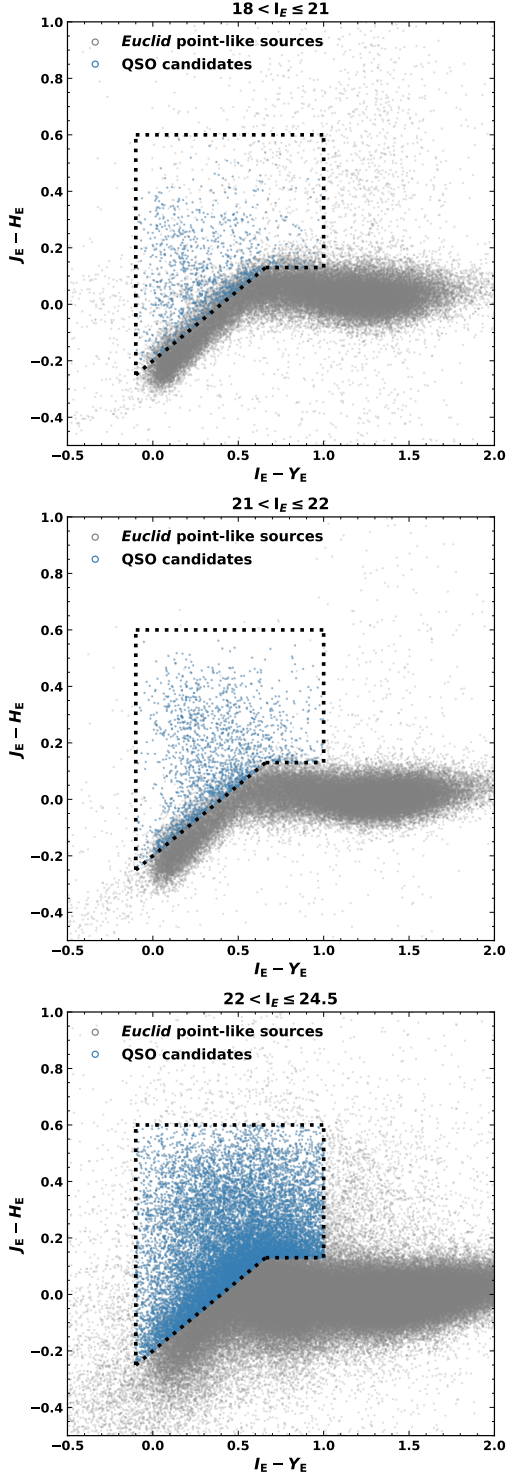


Fig. A.3. Comparison between the number of QSO (blue) candidates per magnitude bin in the EDF-N for the new *Euclid*-only colour cut, $JH_{I_E}Y$. In grey we show all *Euclid* compact sources in the corresponding magnitude bin. As the magnitude bins go from brighter colours (top plot) to faintest (bottom plot) the number of sources and QSO candidates increase.

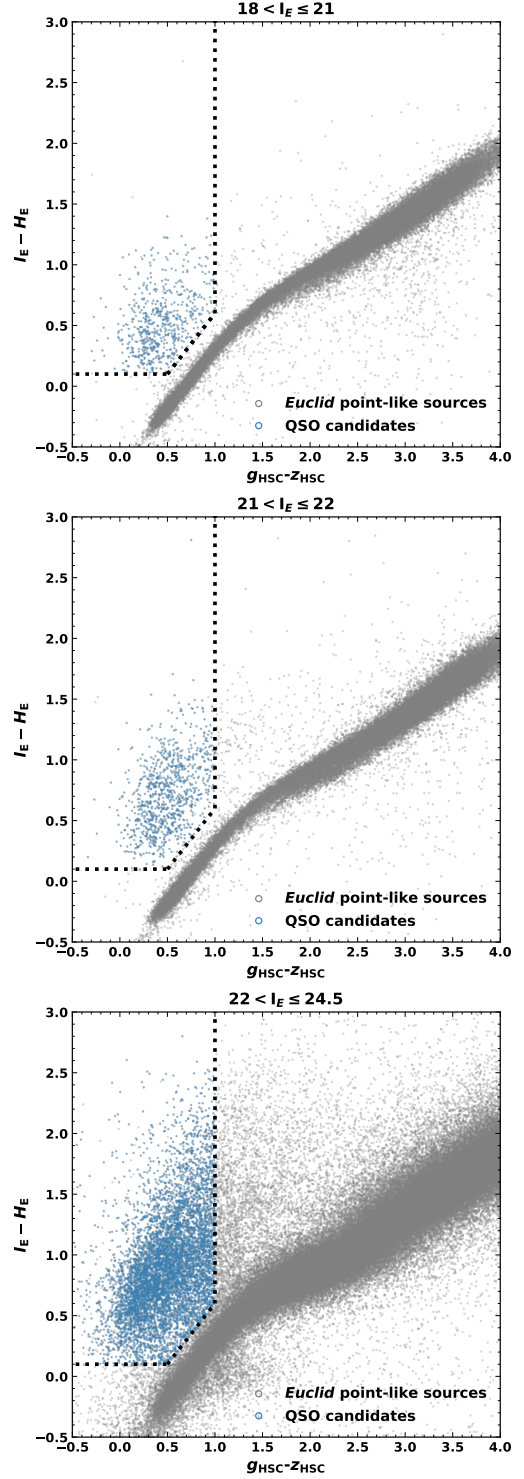


Fig. A.4. Comparison between the number of QSO (blue) candidates per magnitude bin in the EDF-N for the new *Euclid* and ancillary data colour cut, $I_E H_{gz}$. In grey we show all *Euclid* compact sources in the corresponding magnitude bin. As the magnitude bins go from brighter colours (top plot) to faintest (bottom plot) the number of sources and QSO candidates increase.

Appendix B: Comparison between AGN selections

We present the intersection table for the different selection methods investigated in this work, excluding the morphology-based methods due to their distinct methodologies (see Fig. 13). We

limit the depth of every selection to $I_E < 22$ to ensure comparability without bias towards the *Euclid*-based selections, which are the only ones capable of reaching the faintest magnitudes.

Table B.1. EDF-N intersection matrix between the different AGN selection methods investigated in this work, excluding the morphology based ones, and limited to $18 < I_E \leq 22$.

	B24A	B24B	DESI	PRF	GDR3-QSO	JH_{I_EY}	X-ray	I_EH_{gz}	C75	R90	SED fitting
B24A	3123
B24B	1209	1655
DESI	299	433	2886
PRF	26	1	43	2020
GDR3-QSO	298	381	281	15	641
JH_{I_EY}	1091	1000	714	97	490	2873
X-ray	28	41	43	3	24	59	134
I_EH_{gz}	622	927	635	73	457	1400	54	1500
C75	400	563	518	295	437	870	46	769	3187
R90	185	238	256	17	318	369	23	311	685	694	...
SED fitting	294	409	1 505	30	272	635	40	572	488	251	3390

Table B.2. EDF-S intersection matrix between the different AGN selection methods investigated in this work, excluding the morphology based ones, and limited to $18 < I_E \leq 22$.

	B24A	PRF	GDR3-QSO	JH_{I_EY}	X-ray	I_EH_{gz}	C75	R90
B24A	3323
PRF	0	2999
GDR3-QSO	465	0	811
JH_{I_EY}	1607	0	616	3593
X-ray	292	5	315	520	1050
I_EH_{gz}	992	0	557	1857	452	2034
C75	708	700	623	1326	524	1110	5528	...
R90	322	25	457	566	300	456	1123	1139

Table B.3. EDF-F intersection matrix between the different AGN selection methods investigated in this work, excluding the morphology based ones, and limited to $18 < I_E \leq 22$.

	B24A	PRF	GDR3-QSO	JH_{I_EY}	X-ray	I_EH_{gz}	C75	R90
B24A	1366
PRF	0	1535
GDR3-QSO	273	0	510
JH_{I_EY}	720	0	401	1489
X-ray	225	4	213	399	967
I_EH_{gz}	519	0	381	911	341	1011
C75	327	307	396	642	366	549	2917	...
R90	175	19	282	311	180	262	596	603

Appendix C: Column description of AGN catalogue

We list column descriptions for the three EDFs catalogues below.

1. `object_id_euclid`: *Euclid* unique source identifier.
2. `right_ascension_euclid`: *Euclid* source barycentric right ascension coordinate in decimal degrees.
3. `declination_euclid`: *Euclid* source barycenter declination coordinate in decimal degrees.
4. `source_id_gaia`: *Gaia* unique source identifier.
5. `right_ascension_gaia`: *Gaia* barycentric right ascension in ICRS at the reference epoch.
6. `declination_gaia`: *Gaia* barycentric declination in ICRS at the reference epoch.
7. `id_galex`: GALEX merged ID. Extraction identification number.
8. `alpha_j2000_galex`: GALEX right ascension in degrees (0 to 360) in a J2000 reference frame.
9. `delta_j2000_galex`: GALEX declination in degrees (-90 to 90) in a J2000 reference frame.
10. `source_id_allwise`: WISE-AllWISE unique source ID.
11. `right_ascension_allwise`: WISE-AllWISE J2000 right ascension with respect to the 2MASS PSC reference frame from the non-moving source extraction.
12. `declination_allwise`: WISE-AllWISE J2000 declination with respect to the 2MASS PSC reference frame from the non-moving source extraction.
13. `object_id_spitzer`: *Spitzer* unique source identifier.
14. `brick_desi`: DESI brick ID from tractor input.
15. `brickname_desi`: DESI brick name from tractor input.
16. `brick_objid_desi`: DESI imaging surveys OBJID on that brick.
17. `right_ascension_desi`: DESI barycentric right ascension in ICRS.
18. `declination_desi`: DESI barycentric declination in ICRS.
19. `z_desi`: DESI redshift measured by Redrock.
20. `specobjid_sdss`: SDSS object identification number.
21. `right_ascension_sdss`: SDSS right ascension in decimal degrees.
22. `declination_sdss`: SDSS declination in decimal degrees.
23. `coadd_object_id_des`: DES unique identifier for the coadded objects.

24. `right_ascension_des`: DES right ascension, with quantised precision for indexing.
25. `declination_des`: DES declination, with quantised precision for indexing.
26. `good_flags`: cleaning implemented in the work to keep only those sources with ‘good flags’.
27. `bright_vis_mag_bin`: bright I_E magnitude bin: $18 < I_E \leq 21$.
28. `medium_vis_mag_bin`: medium I_E magnitude bin: $21 < I_E \leq 22$.
29. `faint_vis_mag_bin`: faint I_E magnitude bin: $22 < I_E \leq 24.5$.
30. `star_candidate_gaia`: identifier for stars based on *Gaia*’s proper motion and parallax.
31. `star_candidate_prf`: identifier for stars based on PRF probability > 0.7 .
32. `star_candidate_all`: identifier for stars combining *Gaia*’s proper motion and parallax, and PRF probability.
33. `prf_qso_candidate`: identifier for QSOs based on PRF probability > 0.85 in the EDF-N and > 0.95 in the EDF-S and EDF-F.
34. `B24a_qso_candidate`: identifier for QSO candidates based on B24A.
35. `B24b_qso_candidate`: identifier for QSO candidates based on B24B.
36. `C75_agn_candidate`: identifier for AGN candidates based on C75.
37. `R90_agn_candidate`: identifier for AGN candidates based on R90.
38. `GDR3_qso_candidate`: identifier for QSO candidates based on GDR3-QSO.
39. `JH_IeY_qso_candidate`: identifier for QSO candidates based on $JH_{I_E}Y$.
40. `IeH_gz_qso_candidate`: identifier for QSO candidates based on $I_E H_{gz}$.
41. `DESI_broad_qso_candidate`: identifier for QSO candidates based on DESI SPECTYPE and presence of broad emission lines.
42. `DESI_broad_galaxy_candidate`: identifier for AGN candidates based on DESI SPECTYPE and presence of broad emission lines.
43. `DESI_niibpt_agn_candidate`: identifier for AGN candidates based on DESI spectra and N II BPT diagnostic.
44. `DESI_siibpt_agn_candidate`: identifier for AGN candidates based on DESI spectra and S II BPT diagnostic.
45. `DESI_oibpt_agn_candidate`: identifier for AGN candidates based on DESI spectra and O I BPT diagnostic.
46. `DESI_whan_agn_candidate`: identifier for AGN candidates based on DESI spectra and WHAN diagnostic.
47. `DESI_blue_agn_candidate`: identifier for AGN candidates based on DESI spectra and Blue diagnostic.
48. `DESI_kex_agn_candidate`: identifier for AGN candidates based on DESI spectra and KEX diagnostic.
49. `AGN_fraction`: AGN fraction derived from SED fitting.
50. `AGN_fraction_err`: AGN fraction error derived from SED fitting.
51. `AGN_sed_candidate`: AGN candidate based on AGN fraction threshold.